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# Photochemistry on the underside of the mesospheric Na layer

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**Abstract**

Lidar observations of the mesospheric Na layer have revealed considerable diurnal variations particularly on the underside of the main layer where more than an order-of-magnitude increase of Na density has been observed below 80 km after sunrise. In this paper, multi-year Na lidar observations over a full diurnal cycle at Utah State University (USU) (41.8°N, 111.8°W) and a global atmospheric model of Na with 0.5 km vertical resolution in the mesosphere lower thermosphere (WACCM-Na) are utilized to explore the dramatic changes of Na density on the layer underside. Photolysis of the principal reservoir $NaHCO_3$ is shown to be primarily responsible for the increase in Na after sunrise, amplified by the increased rate of reaction of $NaHCO_3$ with atomic H, which is mainly produced from the photolysis of $H_2O$ and the reaction of OH with $O_3$. This finding is further supported by Na lidar observation at USU during the solar eclipse (> 96% totality) event on August 21st, 2017, when a decrease and recovery of the Na density on the underside of the layer were observed. Lastly, the model simulation shows that the Fe density underside around 80 km increases larger and earlier than observed Na changes during sunrise because of the considerably faster photolysis rate of its major reservoir FeOH.

## 1. Introduction

The layer of Na atoms in the upper mesosphere and lower thermosphere (MLT, ~ 80-105 km in altitude), is formed naturally by meteoric ablation along with other metallic layers such as Fe, Mg, Ca and K [Plane et al., 2015]. The climatological variations of this Na layer are known to be mainly controlled by a series of chemical reactions and dynamics, including tides, gravity waves and the mean circulation in the MLT [Plane, 2004, Marsh et al., 2013]. Mesospheric Na atoms are an important tracer in the MLT, where they are observed by resonance fluorescence, either by the lidar technique [Krueger et

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al., 2015] or solar-pumped dayglow from space [Fan et al., 2007]. The Na lidar technique has enabled high temporal and spatial resolution measurements of the mesospheric Na layer since the 1970s [Sandford and Gibson, 1970]. In addition to Na density observations, Na temperature/wind lidars can measure the atmospheric temperature and wind fields over full diurnal cycle by observing Doppler broadening and shifting of the hyperfine structure of one of the Na D lines [Krueger et al., 2015]. Atmospheric observations have been complemented by laboratory kinetic studies of the important reactions which control both the neutral and ion-molecule chemistry of Na in the MLT [Plane, 1999; 2004; Plane et al., 2015], and the development of atmospheric models which satisfactorily reproduce seasonal observations over most latitudes [Plane, 2004; Marsh et al., 2013; Li et al., 2018]

However, less detailed work has been done to investigate the diurnal variations in Na density, especially on the underside of the layer where neutral chemistry dominates. Advances in lidar technology have enabled Na density observations over a full diurnal cycle [Chen et al., 1996; States and Gardner, 1999; Clemesha et al., 2002; Yuan et al., 2012]. Utilizing multi-year observations, Yuan et al. (2012) investigated the diurnal variation and tidal period perturbations of the Na density. These tidal Na perturbations were then used to estimate the tidal vertical wind perturbations [Yuan et al., 2014], showing that, although closely correlated with tidal waves and dominated by tidal wave modulations in the lower thermosphere, the Na diurnal and semidiurnal variations cannot be induced by tidal modulations alone. This is especially the case on the underside of the layer below ~ 90 km, where tidal wave amplitudes are relatively small (see Figure 5a and 5b in Yuan et al., 2012), implying that other mechanisms make a significant contribution to the diurnal variation of the Na layer underside.

Plane et al. (1999) recognized the important role of photochemical reactions in the characterizing the underside of the Na layer, and then measured the photolysis cross sections of several Na-containing molecules – NaO, $NaO_2$, NaOH and $NaHCO_3$, which models show are significant mesospheric reservoir species [Self and Plane, 2002, Marsh et al., 2013]. These cross sections, measured at temperatures appropriate to the MLT, were then used to calculate mesospheric photolysis rates:

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$NaO + h\nu \rightarrow Na + O$ $5.5\times10^{-2}\ s^{-1}$ R1

$NaO_2 + h\nu \rightarrow Na + O_2$ $1.9\times10^{-2}\ s^{-1}$ R2

$NaOH + h\nu \rightarrow Na + OH$ $1.8\times10^{-2}\ s^{-1}$ R3

$NaHCO_3 + h\nu \rightarrow Na + HCO_3$ $1.3\times10^{-4}\ s^{-1}$ R4

These *direct* photochemical reactions release atomic Na during daytime. Furthermore, *indirect* photochemistry also plays a role. The photolysis of $O_2$, $O_3$ and $H_2O$ [Brasseur and Solomon, 2005] lead to the production of H and O (R5 – R8), increasing their concentrations by more than 1 order of magnitude during daytime at an altitude around 80 km [Plane, 2003]. The daily variation of H is further facilitated by the reactions between $HO_2$ and $O/O_3$, which has strong diurnal variations.

$O_2 + h\nu \rightarrow 2O$ R5

$O_3 + h\nu \rightarrow O + O_2$ R6

$H_2O + h\nu \rightarrow H + OH$ R7

$OH + O \rightarrow H + O_2$ R8

$NaHCO_3 + H \rightarrow Na + H_2CO_3$ R9

Thus, H and O reduce these Na compounds listed above to atomic Na [Plane, 2004].

To demonstrate the effect of solar radiation on mesospheric Na layer, including R1-R8, Figure 1 shows the averaged Na density variation in the layer between 75 km and 105 km during a 6-hour period that straddles sunrise (from 3 hours before to 3 hours after sunrise) in the fall season (from August 20 to September 30). The results are based on 50 days of Na lidar observations at USU (41.8°N, 111.8°W) between 2011 and 2016. Figure 1 also includes the ratio profile between the Na density 3 hours after sunrise to that 3 hours before sunrise. The lidar observations clearly show that, while there is an overall Na density increase after sunrise below ~ 92 km, the increase at and below 80 km is much larger than

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closer to the layer peak: the Na density increases by a factor of ~ 6 at 80 km and ~ 40 near 78 km, whereas it is almost uncharged around ~ 95 km. The dramatic oscillation of the ratio below 78 km is due to very low Na density before sunrise (usually below 1 $cm^{-3}$). Note that this ratio calculation for Na density profiles 1 hour before and 1 hour after sunrise generates a similar ratio profile, which demonstrated even larger ratio near 80 km, a factor of ~ 20. This implicates very quick Na density enhancement during the process of sunrise. This analysis therefore provides strong evidence for the impact of photochemistry on the underside of the Na layer.

In this paper we compare the USU Na lidar diurnal cycle observations of mesospheric Na layer during a continuous 7-day campaign in fall 2012 with the Na density variation simulated by NCAR's Whole Atmosphere Community Climate Model with Na chemistry (WACCM-Na) [Marsh et al., 2013] at the USU location during the same period, in order to investigate quantitatively the role of photochemistry on the Na layer. In addition, Fe density variation by the latest WACCM-Fe [Feng et al., 2017] due to photolysis is also discussed to show the distinct feature of the Fe in the underside of the main layer. The Na Lidar measurements made during the solar eclipse on August 21, 2017 in North America are then used as a further robust test of the role of photolysis.

## 2. Instrument and Model description

The USU Na Temperature/Wind lidar system, originally developed at Colorado State University, has been operating at the USU main campus since summer 2010. In addition to Na density observations, neutral temperature and winds are also measured for the mesopause region (~80-110 km) [Krueger et al., 2015]. The lidar return signals can be recorded in 150 m bins in the line-of-sight direction, and saved every minute. Facilitated by a pair of customized Faraday filters deployed at its receiver [Harrell et al., 2009], this advanced lidar system can also reject the sky background significantly during daytime, while receiving the Na echo with minimum loss. This technique provides robust measurements of these important atmospheric parameters under sunlight condition, thereby enabling this investigation of Na photochemistry. In this study, we focus on two sets of Na lidar data: Na density data taken between UT

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Day 271 (September 27) and UT Day 277 (October 3) of 2012 and lidar observations during the solar eclipse on August 21, 2017. The lidar observations presented here are processed with 2-km vertical resolution, and 10-minute and 30-minute temporal resolution for nighttime and daytime data, respectively, to achieve appropriate signal to noise (S/N) for studying the underside of the layer between 75 and 80 km, where the Na density is low (<100 $cm^{-3}$). The lidar observations during the solar eclipse are processed with 10-minute resolution to investigate the potential eclipse-induced perturbations in detail.

WACCM-Na is a global meteoric Na model which satisfactorily reproduces lidar and satellite measurements of the Na layer [e.g., Marsh et al., 2013; Dunker et al., 2015; Plane et al., 2015; Langowski et al., 2017; Dawkins et al., 2016; Feng et al., 2017]. WACCM-Na uses the Community Earth System Model (version 1) framework (e.g., Hurrell et al., 2013), which includes detailed physical processes as described in the Community Atmosphere Model, version 4 (CAM4) [Neale et al., 2012], and has the fully interactive chemistry described in Kinnison et al. (2007). The current configuration for WACCM is based on a finite volume dynamical core [Lin, 2004] for tracer advection. For the present study, we used a specific dynamics (SD) version of WACCM, in which winds and temperatures below 50-60 km are nudged towards NASA's Modern-Era Retrospective Analysis for Research and Applications (MERRA) [Lamarque et al., 2012]. The horizontal resolution is 1.9˚ latitude × 2.5˚ longitude. For this study we performed model experiments using two different vertical resolution: 88 and 144 vertical model levels (termed as lev88 and lev144), both have the same 62 vertical levels from surface to 0.42 Pa (below ~50 km) as MERRA with different vertical resolution above 0.42 Pa ( ie., lev88 which has been used as a standard SD-WACCM gives a coarse height resolution from ~1.9 to ~3.5 km above the upper stratosphere to MLT while lev144 has fine vertical resolution increases from 1.9 km down to around 500 m in the MLT [Merkel et al., 2009; Viehl et al., 2016]. The Na reaction scheme described in Plane et al. (2015) is updated with the results of recent laboratory studies [Gomez-Martin et al., 2016; 2017], and the meteoric input function (MIF) of Na from Cárrillo-Sanchez et al. (2016) is used. Note that the absolute Na MIF used in this paper is the same as in Li et al. (2018) and Plane et al. (2018); i.e., it has been divided by a factor of 5 from the MIF in Cárrillo-Sanchez et al. (2016), to match the observations. In order to contrast the photochemical behavior of the Na layer underside with that of the Fe layer, a WACCM-Fe simulation was also performed. The model output was sampled over USU (41.8°N, 111.8°W) every thirty minutes (the model time step) then interpolated to the same observational period for the available lidar daytime measurements for direct comparisons (note that the modelled nighttime uses the same daytime because the model time step is every 30 minutes while the lidar nighttime measurement is every 10 minutes).

## 3. Comparison of the Na lidar observations with WACCM-Na

The averaged Na density diurnal variation calculated from the intensive 7-day USU Na lidar campaign is presented as the top plot in Figure 2. During the campaign, the time of sunrise in the MLT is around 06:56 local time (LT) based on solar elevation angle (-5° represents sunrise in the MLT); the noon and sunset times are 13:45 LT and 20:35 LT, respectively. The observations reveal strong variations within the layer during the day. Close to the layer peak there are three minima at around 20:00 LT (evening), 04:00 LT (right before dawn) and 16:00 LT (afternoon), where the Na density falls to ~ 3000/cm$^3$. These are separated by two significant maxima near 91 km: the stronger occurs right after sunrise and lasts almost the whole morning with a peak density more than 4400/cm$^3$; the other maximum occurs near 22:00 LT (shortly before midnight), and has a much shorter lifetime (~ 1 hour) with peak density slightly above 4100/cm$^3$. Similar to Figure 1, on the underside of the main layer there is clear evidence of an increase of Na density after sunrise.

Compared with the lidar observations, Figure 2 (middle panel) shows that the relatively coarse resolution WACCM-Na produces a reasonable Na layer in terms of a peak Na density close to 4500 cm$^{-3}$. However, in contrast to the lidar observations, three distinct features are observed: first, the maxima and minima around the layer peak are much less obvious; second, the peak height of the simulated Na layer is near 87 km, about 3-4 km lower than the lidar observations, partly due to a few km lower mesospause in SD-WACCM (Feng et al., 2013) ; third, the absolute value of Na density vertical gradient below the layer peak is much larger than observed. For instance, the modeled Na density decreases from near 4500 $cm^{-3}$ at 87 km to ~ 2000 $cm^{-3}$ around 82 km, while a similar density decrease is observed by the lidar to occur between about 95 km and 82 km. Of course, the second and third differences are probably related. In contrast, Figure 2 (bottom panel) shows that the WACCM-Na high resolution (lev144) output does capture the three-minima at the layer peak during a diurnal cycle, as observed (Figure 2, top panel). Although the Na density near the first minimum (~ 4500 $cm^{-3}$) is higher than observed, the times of the minima, which are close to 20:00 LT, 05:00 LT and 14:00 LT, are in good accord with the lidar observations. However, the Na peak density (> 5500/$cm^3$), along with the overall Na column abundance, is considerably higher than measured by the lidar, and the differences in peak height and vertical density gradient still persist.

In order to examine the density variation on the underside of the layer in greater detail, Figure 3 compares the time-resolved variation of the partial Vertical Column Density (VCD) below 80 km, where the magnitude of variations is the largest, from the 7-day lidar observations in September 2012 and the two WACCM-Na simulations. The variation of solar zenith angle is also plotted. Here, due to the differences of absolute density among the data sets, each VCD is normalized to its maximum so that all can fit in one plot. There is excellent agreement between the measured and modeled rates of change in Na VCD around sunrise and sunset. The VCD reaches a maximum around midday and is then fairly constant in the afternoon. This is because, contrary to the scenario in troposphere, the solar intensity in the mesosphere is pretty constant during the day. During the night, the Na VCD gradually decays to a minimum immediately before sunrise. The WACCM lev144 simulation also captures better the observed rate of decrease of Na immediately after sunset. Note also that the rate of decrease after sunset is faster than the increase following sunrise. This is caused by the rapid decrease in the concentrations of O and H at sunset, compared with their slower photochemical buildup (R5-R8) and the photolysis of Na reservoir species (R1-R4) after sunrise.

An interesting contrast can be made with the behavior of the underside of the mesospheric Fe layer, where considerable density variations due to photolysis have also been observed by Fe lidars within the similar altitude range [Yu et al., 2012; Viehl et al., 2016]. Also included in Figure 3 is the modeled variation of the Fe VCD between 75 and 80 km, using a lev144 simulation with WACCM-Fe, which has been validated against Fe lidar observations [Viehl et al., 2016]. Note that although the rate of decrease of the Fe VCD around sunset is almost identical to that of Na (because both species correlate with the falling O and H concentrations after sunset) the rate of increase at sunrise is significantly faster. The Fe VCD reaches 70% of its daytime maximum within about 1 hour, whereas the Na VCD takes more than 4 hours to reach the same percentage of its maximum.

## 4. Na variation during solar eclipse

During the solar eclipse on August 21 2017, the USU Na lidar conducted a special campaign to observe its potential impact on the MLT. The lidar beam was pointed to the north, 30° off zenith, and operated between 09:45 LT and 15:00 UT. Although this campaign was limited by poor sky conditions in the early morning and afternoon, it was able to cover the complete course of the solar eclipse and observe the MLT at the peak of the eclipse with more than 96% of totality at 11:34 LT. To our knowledge, these are the first lidar observations of the neutral Na in the MLT during an eclipse.

Figure 4 shows the lidar-observed variations of mesospheric Na density throughout the eclipse with10-minute temporal resolution. The averaged return signal between 200 and 220 km altitude per lidar Line-of-Sight (LOS) binning range (150 m) is treated as the sky background, which is also shown in Figure 4. The background variation indicates that at the USU location the eclipse began at 10:25 LT, peaked at 11:34 LT, and ended just before 13:00 LT. The high background before 10:00 LT was due to hazy sky condition in the early morning, similar to after 15:00 when it became cloudy. During the course of this event, the mesospheric Na layer weakened with decreasing peak density. In particular,  Na density variation was more evident below 85 km. As Figure 4 shows, before the eclipse the constant density lines on the underside of the layer were moving downwards (i.e. increasing density at each altitude). As the eclipse unfolded, these constant density lines started to move towards higher altitudes (the density decreased at each altitude). During the recovering phase of the eclipse, the Na density began to increase again. By ~ 13:00 LT, the Na layer was fully recovered, and no significant change was observed on the underside of the layer. For example, the density line of 450 $cm^{-3}$ was near 80 km right before the eclipse started. It moved upward to near 81.5 km at the culmination of the event, before went back and stayed near 80 km at the end of the event. Similar behavior can be seen in almost all constant density line below 85 km. Further calculation of the lidar measured underside Na VCD (75-85 km) shows that it decreased as much as about 40% between 10:25 LT and 11:35 LT.

The simultaneous temperature and horizontal wind measurements during the eclipse, however, do not reveal apparent variations that can be associated with the event (not shown). For instance, the measured temperature change is within the daytime lidar measurement uncertainty (~ 5K with 4 km smoothing). Such a small temperature change during the solar eclipse is expected when considering the general energy budget in the MLT. When the short-wave heating that dominates the daytime budget (mainly exothermic heating from atomic O recombination is turned off (Brasseur and Solomon, 2005)), infrared cooling due to $CO_2$ emission would lead to net cooling in the mesopause region. However, the magnitude of this cooling is only about 1 K/hour (Roble, 1995). Thus, for a solar eclipse that only lasts for two hours with just a few minutes of totality, a noticeable temperature change should not occur. This result is consistent with a recent simulation of the eclipse using the WACCM-X model, which concluded that the temperature variation in the mesosphere would have been no more than 4 K [McInerney et al., 2018]. Furthermore, the variation of temperature within this range will not have a significant impact on the Na reaction kinetics and hence the Na atom density [Plane et al., 1999].

## 5. Discussion

The major reservoir for Na on the underside of the layer is $NaHCO_3$, which forms by three steps from Na atoms: oxidation of Na by $O_3$ to NaO; reaction with $H_2O$ or $H_2$ to form NaOH; and recombination of NaOH with $CO_2$ [Plane et al., 2015; Gomez-Martin et al., 2017]. $NaHCO_3$ is converted back to Na either by photolysis (reaction R4), or by reaction with atom H in R9. The rate coefficient for this reaction, R9, is $k_9 = 1.84 \times 10^{-13}\ T^{0.78}\ \exp(-1014/T)$ cm$^3$ molecule$^{-1}$ s$^{-1}$ [Cox et al., 2001], which is $7.1 \times 10^{-14}$ cm$^3$ molecule$^{-1}$ s$^{-1}$. The typical daytime H concentration between 75 and 80 km is around $5 \times 10^7$ cm$^{-3}$ [Plane et al., 2015], so the first-order rate of this reaction is ~ $4 \times 10^{-6}$ s$^{-1}$, which is about 40 times slower than R4. Thus, photolysis of $NaHCO_3$, which has built up during the night, is responsible for ~98% of the increase in the Na VCD after sunrise (Figure 3). The excellent agreement between the laboratory-measured photolysis rate of $NaHCO_3$ [Self and Plane, 2002], and the observed increase of the Na VCD is strong evidence that $NaHCO_3$ is indeed the major Na reservoir on the underside of the Na layer. Indeed, as shown in Figure 5, the WACCM_Na simulated variations of Na, H and O, between local midnight and noon, are highly correlated in the upper mesosphere in the same 144-level run. Within the same region, the $NaHCO_3$ density decreases above 78 km after sunrise due to photolysis and the increases in H and O. Below this altitude, the change in $NaHCO_3$ is much smaller because the concentrations of O and H, despite increasing after sunrise, are too small to prevent any Na produced being rapidly converted back to $NaHCO_3$. This is because the initial oxidation step – the recombination reaction of Na with $O_2$ (which is a pressure-dependent reaction) – becomes very fast below 78 km.

Based on the above discussion, a decrease of Na on the underside of the layer would be expected during the eclipse, because of the reduction in the photolysis rates (R1-R4) and atomic O and H. As shown in Figure 4, a decrease in the Na density below 85 km was indeed observed by the lidar during this period. However, the decrease in Na is relatively small when compared with the natural variability measured with the USU Na lidar during the morning hours (50 days of lidar data between August 20 and September 30 in 2011-2016), which can be as much as 60% between 75 and 85 km and is mostly driven by atmospheric gravity wave modulations [Shelton et al., 1980]. The effect of the eclipse was modeled by driving a 1D model of the Na layer (0.5 km vertical resolution, 5 min time resolution) [Plane, 2004] with the background atmospheric species ($O_3$, O, H etc.) from the WACCM-X simulation of the eclipse [McInerney et al., 2018]. As shown in Figure 6, the modeled decrease in Na around 80 km is much less than the diurnal change in the layer because the period where photolysis is significantly reduced during the eclipse is too short, and hence within the natural variability of the layer underside.

In terms of the mesospheric Fe layer, the major Fe reservoir in this region is most likely FeOH [Self and Plane 2003; Plane 2004; Plane et al., 2015]. Similar to Na, the dominant Fe production processes within the bottom half of the layer involve reaction with H and photolysis:

$FeOH + H \rightarrow Fe + H_2O$ $3.1 \times 10^{-10} \exp(-1264/T)$ cm$^3$ molecule$^{-1}$ s$^{-1}$ (R10)

$FeOH + h\nu \rightarrow Fe + OH$ $6 \times 10^{-3}$ s$^{-1}$ (R11)

The rate coefficients of both these reactions are considerably higher than those of the analogous Na reactions, particularly R11 which is faster than R4 by more than two orders of magnitude [Viehl et al., 2016]. It is this feature which controls the more rapid appearance of Fe around 80 km after sunrise, as shown in Figure 4. Note that this rapid increase has been previously observed by lidar [Viehl et al., 2016; Yu et al., 2012].

## 6. Conclusions

Observations of the full diurnal cycle of the underside of the mesospheric Na layer reveal substantial changes in Na density near and below 80 km, with more than an order-of-magnitude increase after sunrise, while the change of Na density above 90 km during the same process is relatively slow. In this study we show that this diurnal variation is largely driven by the photochemistry of the major reservoir species $NaHCO_3$. This result is established by demonstrating reasonable agreement between USU lidar observations of the Na layer below 80 km, and a whole atmosphere chemistry-climate model which includes a comprehensive Na chemistry module (WACCM-Na). Indirect photochemistry, where atomic H and O are produced by the photolysis of $O_3$, $O_2$ and $H_2O$, and these atoms then reduce Na compounds ($NaHCO_3$, NaOH, NaO and $NaO_2$) back to Na, also plays an important role in the diurnal variability. The more rapid increase of atomic Fe after sunrise, which has been observed in several locations [Viehl et al., 2016; Yu et al., 2012], is consistent with the much faster rate of photolysis of FeOH compared with $NaHCO_3$. Lidar observations made during the solar eclipse on August 21, 2017 (at a location with 96%

totality) did not reveal significant changes in either temperature or Na density that were larger than the natural variability around 80 km. This is consistent with a recent study using WACCM-X [McInerney et al., 2018] and the Na model results presented here.

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**Figure Captions:**

Figure 1. The Na density ($cm^{-3}$) variation between 75 km and 105 km in MLT during sunrise between August 20$^{th}$ and September 30$^{th}$ (contour plot) in 2011-2016. Zero hour marks time of the sunrise at the mesopause (bottom abscissa). The solid black profile is the ratio between the Na density 3 hours after sunrise to that 3 hours before sunrise (top abscissa).

Figure 2. The averaged lidar measured Na density variation during the 7-day Na lidar campaign between September 27 and October 3, 2012 (top); the Na density variations at USU location during the same time frame simulated by WACCM_Na 88-level (middle) and 144-level (bottom).

Figure 3. The variations of Na VCD (75-80 km) measured the USU Na lidar during the 7-day Na lidar campaign (asterisks), simulated by WACCM_Na 88-level run (orange dotted line) and 144-level run (orange dashed line) and solar zenith angle (black long-dashed line), along with Fe VCD (75-80 km) simulated by WACCM_Fe 144-level run (blue solid line).

Figure 4. The mesospheric Na density variation during the solar eclipse on August 2, 2017, observed by the Na lidar at Utah State University. The yellow solid line represents the lidar-detected sky background. The unit for the lidar background measurement is #photon/bin/minute.

Figure 5. The variations of Na (a), NaHCO3(b), H(c) and O(d) in the underside of the mesospheric Na layer, simulated by WACCM_Na_144-level run.

Figure 6. A 1D model simulation of the Na layer variation during the solar eclipse between 18:00 LT on August 20 and 18:00 LT on August 21 in 2017. The solar eclipse at the USU Na lidar location peaked at 11:34 on August 21 (marked by the solid arrow). The background atmospheric species ($O_3$, O, H etc.) are based on the outputs of WACCM-X eclipse simulation.

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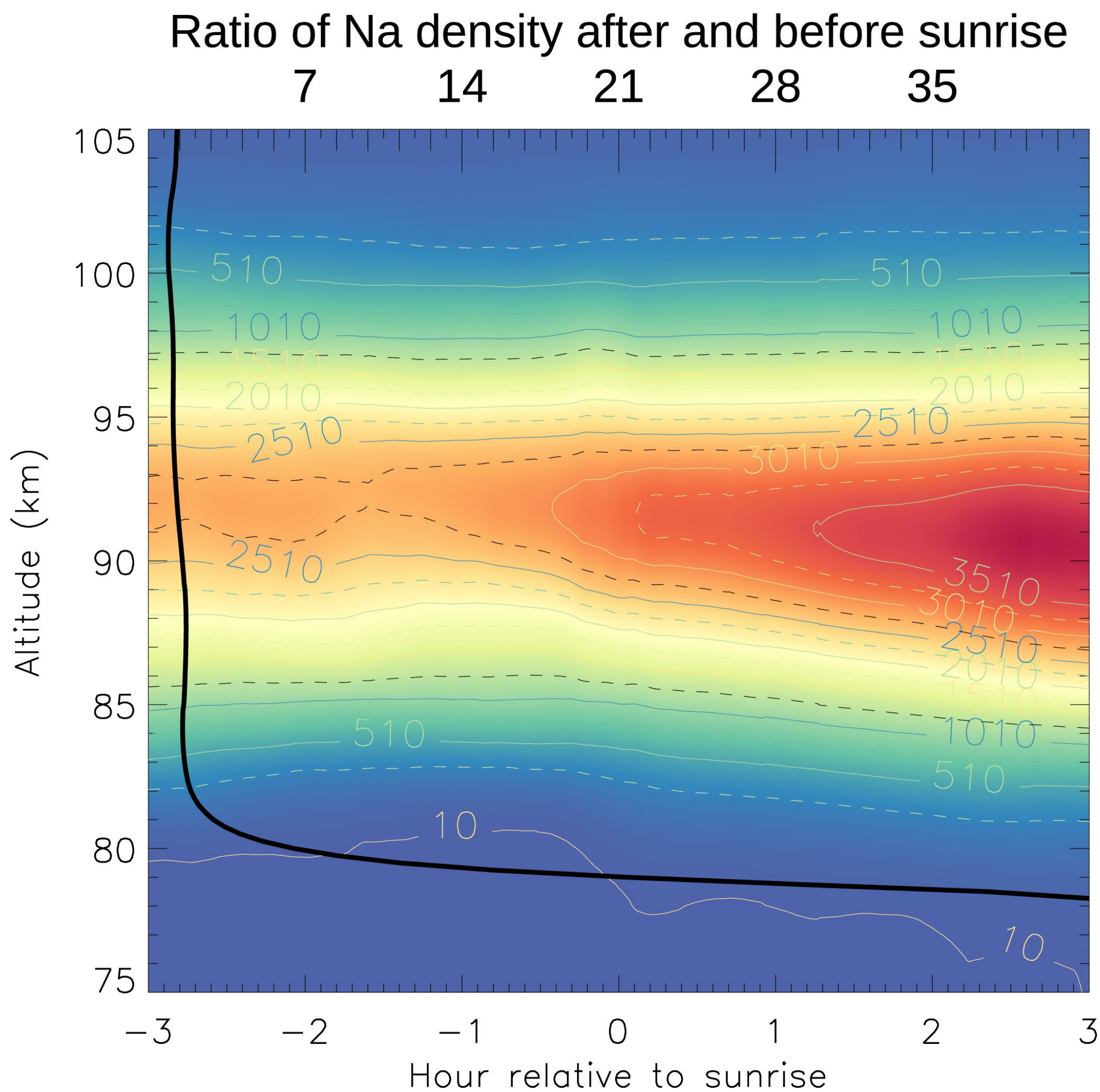


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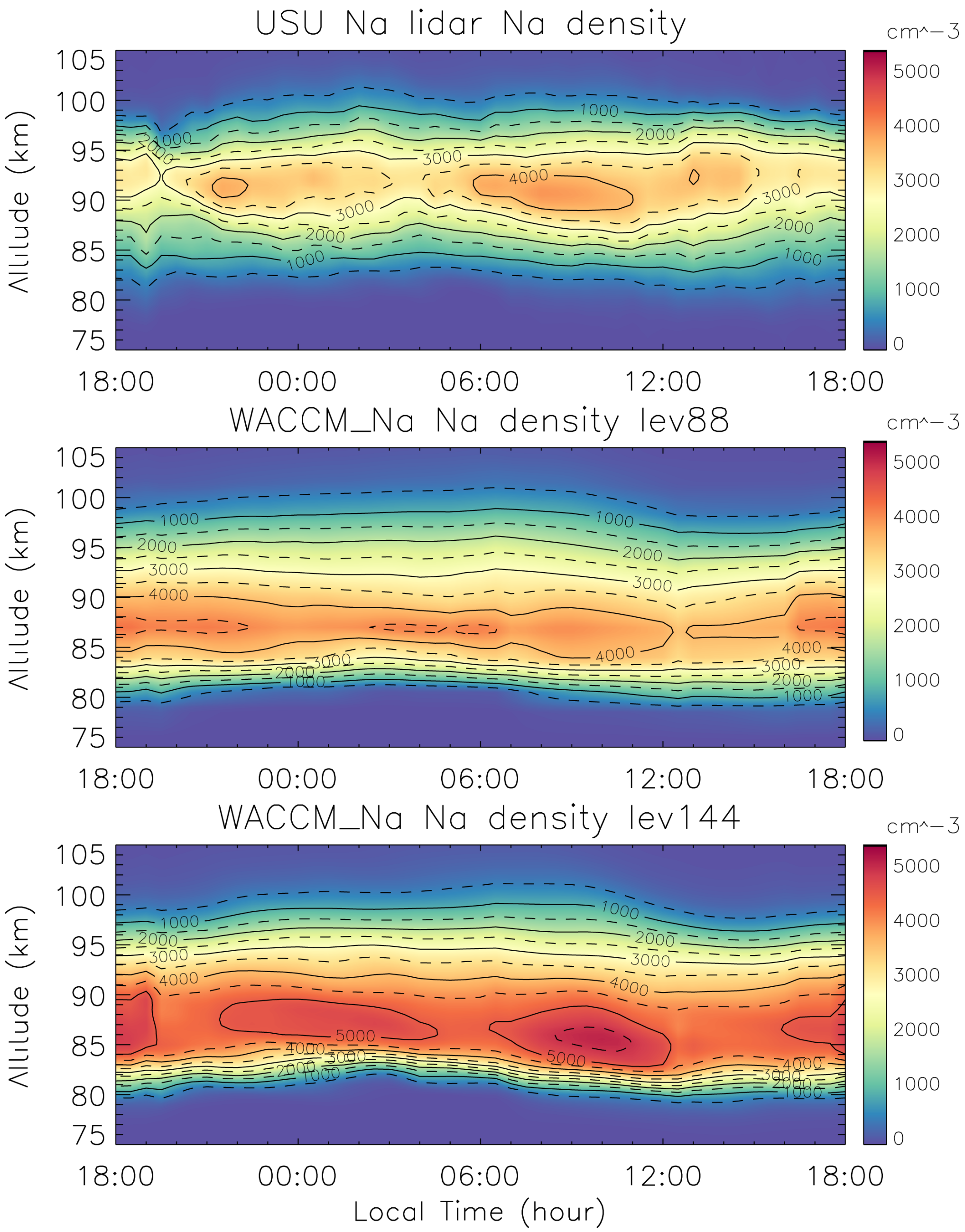


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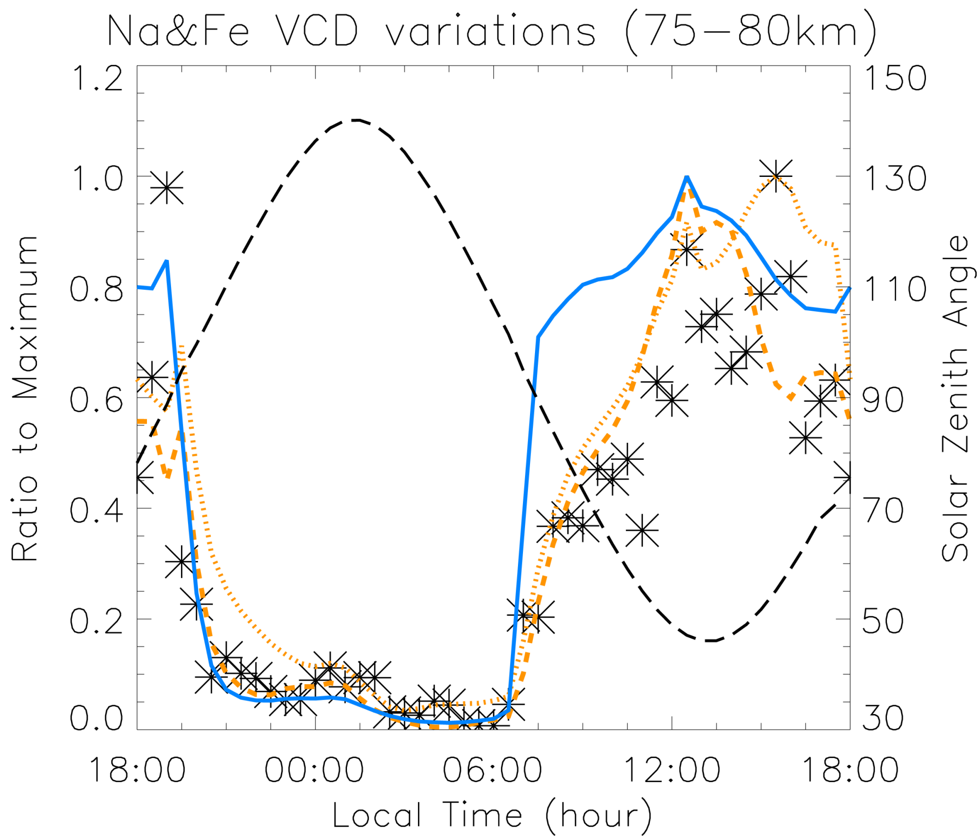


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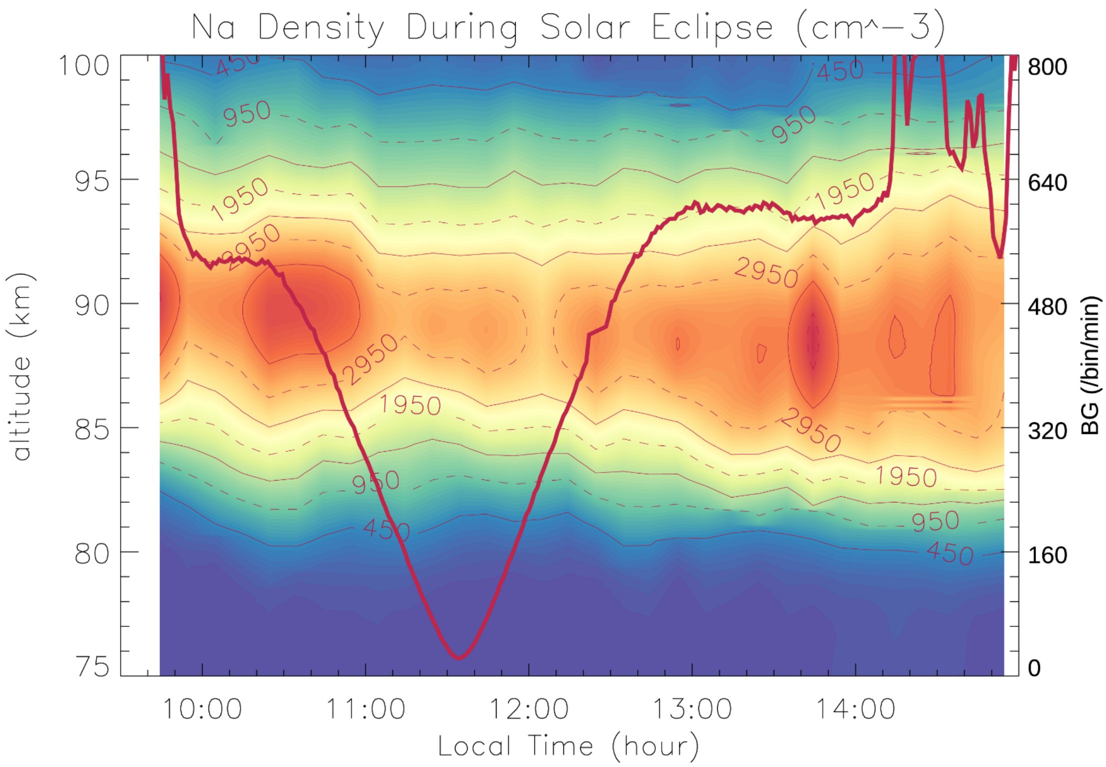


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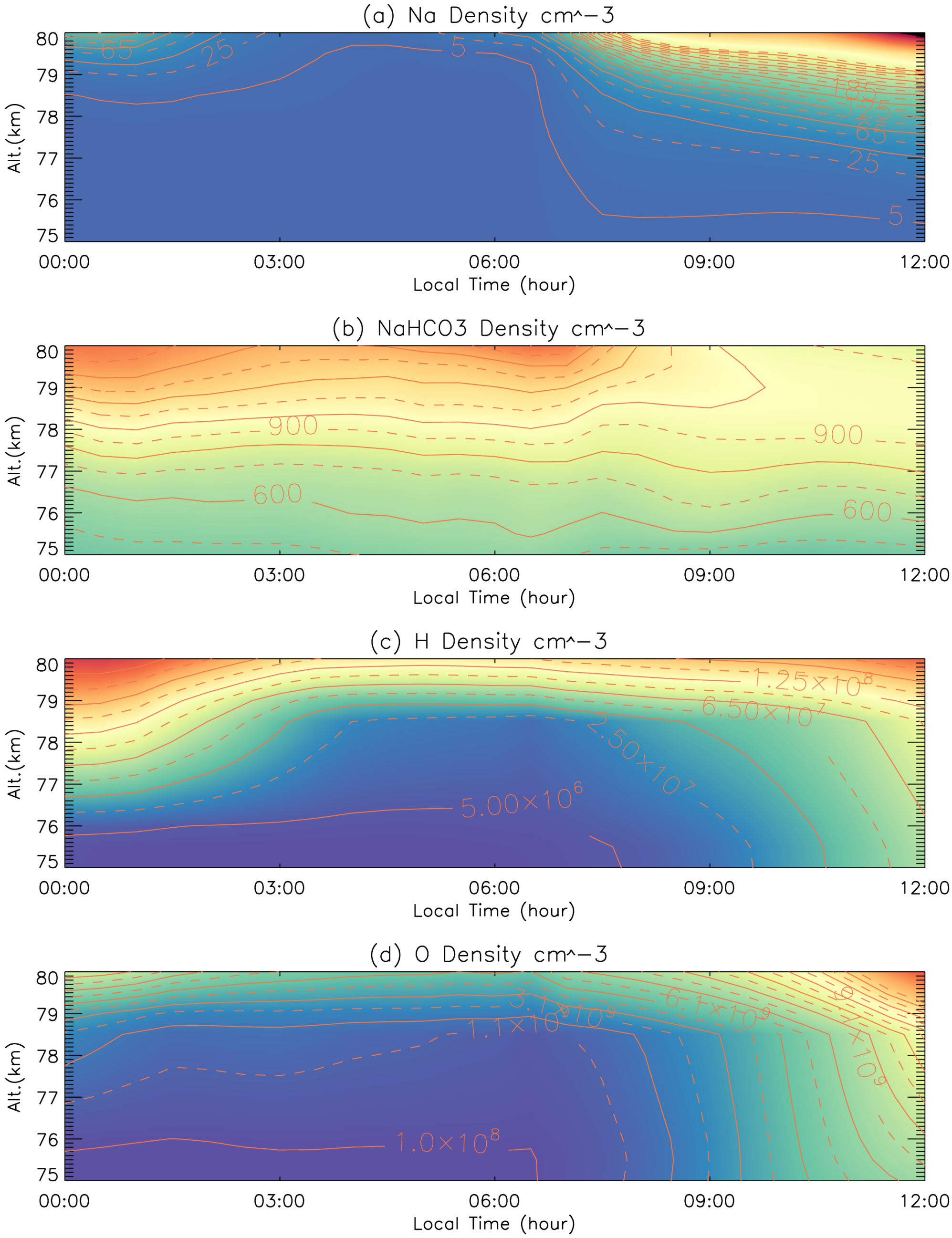


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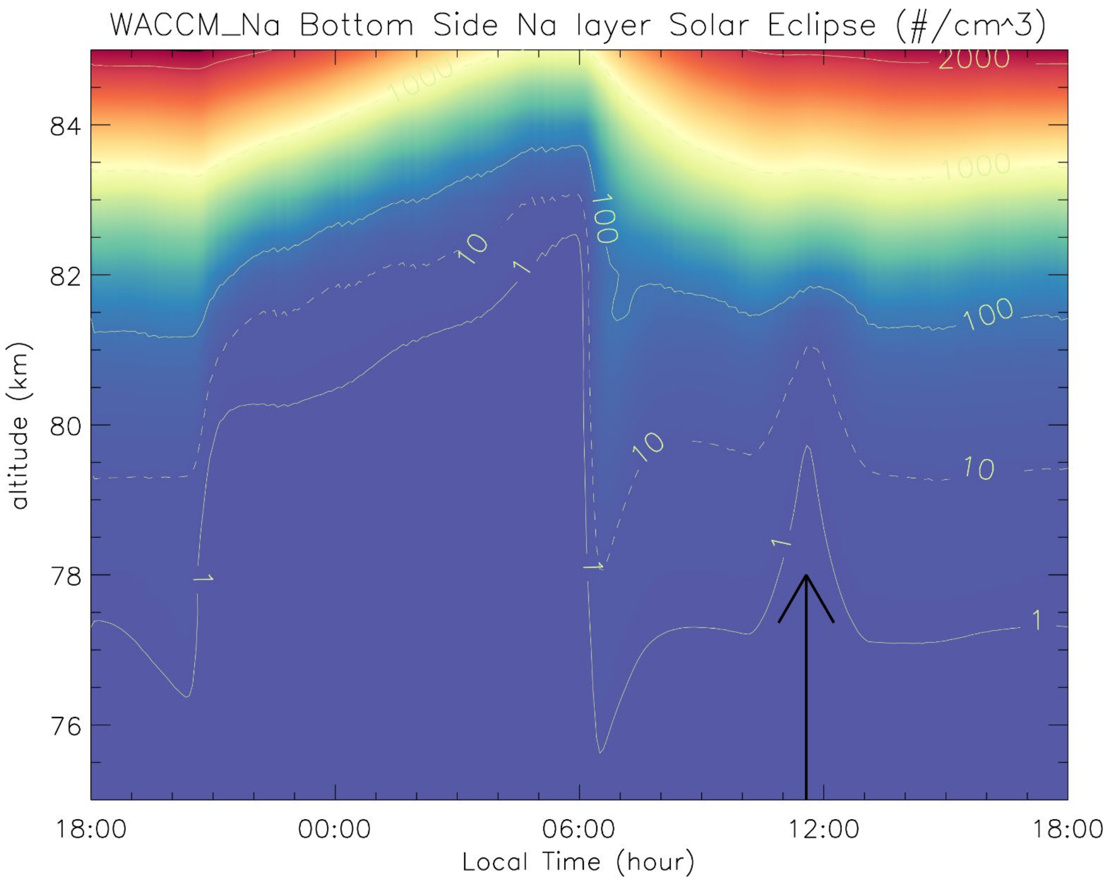


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