

Spatio-temporal observations of tertiary ozone maximum

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Abstract

We present spatio-temporal distributions of tertiary ozone maximum (TOM), based on GOMOS (Global Ozone Monitoring by Occultation of Stars) ozone measurements in 2002-2006. The tertiary ozone maximum is typically observed in the high-latitude winter mesosphere at altitude ~ 72 km. Although the explanation for this phenomenon has been found recently - low concentrations of odd-hydrogen cause the subsequent decrease in odd-oxygen losses - models have had significant deviations from existing observations until recently. Good coverage of polar night regions by GOMOS data has allowed for the first time obtaining spatial and temporal observational distributions of night-time ozone mixing ratio in the mesosphere.

The distributions obtained from GOMOS data have specific features, which are variable from year to year. In particular, due to a long lifetime of ozone in polar night conditions, the downward transport of polar air by the meridional circulation is clearly observed in the

tertiary ozone maximum time series. Although the maximum tertiary ozone mixing ratio is achieved close to the polar night terminator (as predicted by the theory), TOM can be observed also at very high latitudes, not only in the beginning and at the end, but also in the middle of winter. We have compared the observational spatio-temporal distributions of tertiary ozone maximum with that obtained using WACCM (Whole Atmosphere Community Climate Model) and found that the specific features are reproduced satisfactorily by the model.

Since ozone in the mesosphere is very sensitive to HO_x concentrations, energetic particle precipitation can significantly modify the shape of the ozone profiles. In particular, GOMOS observations have shown that the tertiary ozone maximum was temporarily destroyed during the January 2005 and December 2006 solar proton events as a result of the HO_x enhancement from the increased ionization.

1 Introduction

Catalytic reaction chains are important for the ozone budget in the middle atmosphere (e.g. Grenfell et al., 2006). In the mesosphere, the most important catalyst affecting ozone is odd hydrogen ($\text{HO}_x = \text{H} + \text{OH} + \text{HO}_2$ etc.). Features such as the secondary and tertiary maxima in ozone profiles are caused by changes in HO_x production with altitude.

The tertiary maximum in ozone mixing ratio profiles reported first by Marsh et al. (2001) is observed at ~ 72 km altitude near the polar night terminator. Modeling results have shown that the maximum is caused by low concentrations of odd-hydrogen and the subsequent decrease in odd-oxygen losses through catalytic cycles involving hydroxyl. The first modelling results (Marsh et al., 2001) reproduced well the position of tertiary maximum, but the predicted peak amplitude (~ 7 ppmv) was more than twice that shown by MLS and CRISTA observations (~ 3 ppmv). Investigations performed in (Hartogh et al., 2004) have shown that the strong underestimation of the peak concentrations due to insufficient vertical resolution of measurements is unlikely. The ozone measurements by the GOMOS (Global Ozone Monitoring by Occultation of Stars) instrument on board the Envisat satellite, which have a very good vertical resolution and accuracy (see Sect. 2), have confirmed the magnitude of tertiary ozone maximum of 2-4 ppmv (Sofieva et al., 2004b).

Hartogh et al. (2004) pointed out that the distribution of odd oxygen during polar night conditions depends essentially on the complex history before the polar night started, on

vertical transport from the thermosphere and on horizontal transport from the domain outside the polar night area. By using the combined dynamical and chemical transport model COMMA-IAP, the authors got the tertiary ozone mixing ratio close to the observations by the ground-based millimeter radiometer at ALOMAR, and presented detailed spatio-temporal distributions of mesospheric ozone.

To our knowledge, a description of the mean structure and seasonal variability of the observed tertiary ozone maximum has not been published thus far. In this paper, we present the first spatio-temporal distributions of tertiary ozone maximum obtained from ozone measurements by GOMOS. We compare the experimental TOM distributions with that simulated by the WACCM model. We discuss the influence of dynamics and energetic particle precipitation on variability of the tertiary ozone maximum.

2 Ozone measurements by GOMOS on Envisat

GOMOS (<http://envisat.esa.int/instruments/gomos>; Kyrölä et al., 2004; Bertaux et al., 2004) on board the Envisat satellite is the first operational instrument that uses starlight for monitoring the atmospheric composition. The GOMOS spectrometers measure the stellar spectrum continuously as a star sets behind the Earth limb. Vertical profiles of ozone, NO₂, NO₃, and aerosols extinction are retrieved from the UV-Visible spectrometer measurements. Since the launch in March 2002, GOMOS has performed more than 600,000 occultations.

The basis for the geophysical data retrieval from GOMOS measurements is the transmission spectra, which are obtained by dividing the spectra measured at different tangent altitudes by the reference spectrum measured above the atmosphere. In the GOMOS data processing, the inversion is split into two parts: the spectral inversion part and the vertical inversion part (Kyrölä et al., 1993). In the spectral inversion, horizontal column densities are retrieved from the atmospheric transmission data from which refractive effects and scintillation have been removed. In the vertical inversion, vertical profiles are reconstructed from the horizontal column densities. The inversion is stabilized by Tikhonov-type regularization according to the target resolution (Sofieva et al., 2004a; Tamminen et al., 2004), which makes the vertical resolution practically independent of angles between the orbital plane and the direction to the star. In the middle and upper mesosphere, the vertical resolution (including the smoothing properties of the inversion) of GOMOS ozone profiles is about 3 km.

For ozone, the valid altitude range is $\sim 10\text{-}100$ km. The accuracy of the retrieval depends on stellar magnitude and spectral class. Only occultations of hot stars, which emit sufficiently at ultraviolet wavelengths, are able to provide information about ozone in the mesosphere and the thermosphere (Kyrölä et al., 2006). For the analysis of the tertiary ozone maximum, we selected dark limb occultations (solar zenith angle $\geq 107^\circ$ at the tangent point, and $\geq 90^\circ$ at the satellite) of stars with effective temperature $\geq 6000\text{K}$. The accuracy of individual selected profiles at ~ 70 km is $\sim 1.5\text{-}7\%$, depending on stellar brightness.

In our analysis, we used GOMOS ozone data from 2002-2006. The spatio-temporal coverage of the selected nighttime GOMOS observations (131500 occultations) is shown in Figure 1. It is characterized by a good coverage of winter poles, while summer poles are not covered due to the absence of dark limb conditions. During May-June 2003 and February-July 2005, GOMOS suffered from pointing system malfunction, thus data are missing in these periods. Reduced azimuth angle range since August 2005 results in a decreased amount of GOMOS data since that time. The successive occultations of each star are nearly uniformly distributed in the longitudinal direction (maximum 14 occultations per day), and they are carried out approximately at the same local time. Therefore, being averaged over longitude, GOMOS measurements are a good representation of zonal mean data.

GOMOS data are very convenient for interannual comparisons because occultations of the same stars are carried out approximately at the same locations and time in different years. Ozone number density is retrieved from GOMOS measurements. For calculation of mixing ratio, we used combined air density profiles based on ECMWF data in the stratosphere and on MSIS90 model (Hedin, 1990) in the mesosphere, at occultation locations. [A procedure that ensures validity of the hydrostatic equation for the resulting profiles was applied.](#)

3 Brief description of the WACCM model

The Whole Atmosphere Community Climate Model (WACCM) was developed under the leadership of the National Center for Atmospheric Research (NCAR) for investigation of interactions over the depth of the atmosphere. WACCM is a global model that extends from the ground to 4.5×10^{-6} hPa, approximately 145 km. The model is an extension of the Community Atmosphere Model and uses the physical parameterizations from it (see Collins et al., 2004). The dynamical core of WACCM is based on the finite volume method of Lin

(2004). The model includes fully interactive chemistry with 51 neutral species, 5 ions and electrons. It incorporates most of the physical and chemical mechanisms believed to be important for determining the dynamical and chemical structure of the middle atmosphere, including the mesosphere and lower thermosphere. Extensions to physics include parameterizations of electron precipitation in the auroral oval, non-orographic gravity wave drag and molecular diffusion. A full description of the model dynamics, chemistry, radiation, and physical parameterizations is given in Garcia et al., (2007) and Marsh et al. (2007).

The WACCM simulations were carried out at $4^\circ \times 5^\circ$ (latitude \times longitude) resolution. The results of three ensemble simulations described in (Garcia et al., 2007) are used for the analysis presented in this paper; we present the mean distributions over the three ensemble realizations.

4 Spatio-temporal distributions of tertiary ozone maximum

The explanation for the tertiary ozone maximum is that it occurs as a result of O_x loss by catalytic cycles of OH and HO_2 that is decreasing because HO_x production from photolysis of water vapor decreases with reducing UV radiation at high latitudes in winter (Marsh et al., 2001; Hartogh et al., 2004). Thus, the maximum mixing ratio should be achieved close to the polar night terminator. Figure 2 (top) shows ozone mixing ratio at 72 km obtained from the GOMOS data in the period Sep 2002 - Dec 2006. For producing the GOMOS distribution, data were first averaged in $5^\circ \times 10$ days (latitude \times time) bins, and then three-point smoothing (both in latitude and time) was applied to the data field. Although the maximum mixing ratio is achieved close to the polar night terminator (as predicted by the theory), the tertiary ozone maximum is observed also at very high latitudes, not only in the beginning and at the end, but also in the middle of winter (that was not predicted by the modelling results presented in (Hartogh et al., 2004), Figure 11 therein). However, WACCM reproduces very well the evolution of latitudinal distribution of ozone mixing ratio at ~ 72 km (Figure 2, bottom), including also specific features of this distribution (temporal and latitudinal averaging is nearly the same for both experimental and model distributions). The only difference in these distributions is that WACCM predicts slightly larger ozone mixing ratios (up to 5 ppmv) than that observed by GOMOS (not exceeding 4 ppmv).

Figure 3 shows altitude-latitude dependence of GOMOS ozone mixing ratios in the longitudinal band from 30W to 30 E averaged over the period 15 December-15 January, for four NH winters. The latitudinal grid used in Figure 3 is not uniform; it is spaced according to data availability (see Figure 1). At high latitudes, the width of latitudinal bins is 5°. The peak mixing ratio is located at altitudes 65-75km, and usually its latitudinal position is close to the polar night terminator. Unlike the modelling results by Hartogh et al. (2004), the tertiary ozone maximum is not clearly isolated from the high latitudes as predicted by Fig.9 in (Hartogh et al., 2004). In addition, (Hartogh et al., 2004) predicts large ozone values around the pole at upper altitudes (down to 75 km) indicating strong downward transport of atomic oxygen from the thermosphere, which is not confirmed by the GOMOS data.

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Figure 4 shows the analogous latitude-altitude sections obtained from WACCM simulations. The midnight model data for 0° were averaged over three ensemble simulations and over the same time period 15 December - 15 January, as the GOMOS data. The distributions look similar, but the model ozone mixing ratios are ~50% higher than in the GOMOS data. As seen in Fig.3, ozone mixing ratio exhibit rather large variations from year to year. Since WACCM is a free-running climate model, the winds generated within the model are unlikely to match the observed winds on any one day. This difference in meteorology could contribute to the apparent discrepancy between model and observations. The overestimate of the magnitude of the TOM by WACCM is significantly smaller than it has been seen in the ROSE model (Marsh et al, 2001). Since this phenomenon involves hydrogen chemistry and occurs only at high solar zenith angles, it suggests that the model overestimates could stem from inaccuracies in the simulated water distribution in the winter mesosphere or problems with the photolysis rate of water at near-polar night conditions.

According to Figures 3 and 4, the tertiary ozone maximum can extend to the pole. This extension varies from year to year. It seems that the appearance of tertiary ozone maximum close to the pole is related to the intensity of horizontal mixing at high latitudes (see also below).

It is interesting that the altitude position of the tertiary ozone maximum shifts downward with moving closer to the pole. This might be related to the downward transport in the polar night area and can illustrate the peculiarities of the meridional circulation.

Deleted: This disagreement is unlikely explained by the remaining differences in data averaging (GOMOS zonal mean versus WACCM data for 0° longitude averaged over the three ensemble realizations): from GOMOS measurement, we selected data around 0° longitude, and got the distribution very similar to that shown in Figure 3. The diurnal variations cannot also explain the observed differences in TOM magnitudes: GOMOS measurements at high latitudes in winter were carried out between 10 pm and 2 am. ¶

5 Dynamical aspects

In order to follow the evolution of tertiary ozone during polar night, we selected two latitudinal bands in NH: 70°-76° N and 80°-90° N. These latitudinal bands have a very good daily coverage by occultations of the brightest stars (Sirius, Rigel, Procyon), which allow very accurate ozone measurements (uncertainty of individual profiles is ~1.5-2 % at altitudes ~70 km).

The time series of ozone mixing ratio in the mesosphere in two latitudinal bands, for four Northern Hemisphere winters are shown in Figure 5. Based on the modelling results by Hartogh et al. (2004 Figs. 5 and 6), one would expect that the seasonal variations of tertiary ozone maximum at these latitudes consist of two isolated peaks in the beginning (October-November) and at the end (February-March) of winter, according to the variation of the polar night terminator. However, the seasonal variations of ozone mixing ratio in Figure 5 do not show such behavior. As seen in Figure 5, the tertiary ozone maximum is observed usually throughout the winter, but its magnitude is subject to significant variations.

During sudden stratospheric warmings, the tertiary ozone mixing ratio can decrease significantly, as during the major stratospheric warming in December 2003 (Manney et al., 2005), or even the tertiary ozone maximum can be completely destroyed, as during the stratospheric warming in late January 2006, as seen in Figure 5. Due to a long lifetime of ozone in polar night conditions, the downward transport of the polar air by the meridional circulation can be observed in the tertiary ozone time series. In particular, two very strong downward transport events - in January 2004 discussed in (Hauchecorne et al., 2007) and in February-March 2006 discussed in (Randall et al., 2006) - are clearly observed in Figure 5. It is interesting that strong downward transport is observed after sudden stratospheric warmings.

Comparing Figure 5 and Figure 3, one might suggest that there exists a relationship between the isolation of the tertiary ozone maximum from the pole and intensity of downward transport. Since the data shown in Figure 3 are averaged over the period 15 Dec- 15 Jan, only one subplot in Figure 3, for winter 2003-2004, includes the period of strong downward transport. The tertiary ozone maximum is more isolated from the pole for 2003-2004 than for other years. It seems that cases with strong downward transport also have a decrease in horizontal mixing, thus resulting in a better isolation of tertiary ozone maximum from the pole, and vice versa.

6 Influence of energetic particles precipitation

Energetic particle precipitation events, such as Solar Proton Events (SPE), produce HO_x in the high latitude middle atmosphere, including the polar night terminator region where the tertiary ozone maximum is formed. Increased ionization caused by particle precipitation leads to HO_x production via ion chemical reactions (e.g., Solomon et al., 1981). Theoretical predictions of significant HO_x production and the subsequent ozone loss have been verified by the observations of MLS/Aura and GOMOS/Envisat instruments (Verronen et al., 2006). In the terminator area the HO_x concentration is relatively low, therefore moderate SPE production can compensate for the reduced natural HO_x production in the terminator area and can even exceed it. Seppälä et al. (2006) have shown that the tertiary ozone maximum disappeared during the January 2005 SPE. Figure 6 presents another example/confirmation of TOM sensitivity to particle forcing, the destruction of TOM by the moderate SPE of December 2006. Larger values ~ 2 ppmv are seen at 65-70 km until Dec 7 when the SPE began. After the onset of the SPE, the ozone mixing ratio drops below 0.5 ppmv. The data show TOM recovery on Dec 16 when the SPE particle forcing declines to the quiet-time level. Because of the short photochemical lifetime of HO_x , low ozone values are observed only when the particle fluxes are elevated. As the forcing stops, the HO_x production also stops and the HO_x amount returns to normal level within a day, thus resulting in ozone recovery. SPE cause pronounced changes in TOM, because they are sporadic and typically last only for several days. On the other hand, electron precipitation in the auroral oval region also affects TOM. Electron precipitation is more continuous but of lower intensity compared to SPE. For example, in Figure 5 the largest ozone mixing ratios are observed in Jan-Feb 2006 when energetic particle forcing was lowest in 2002-2006 as indicated by geomagnetic activity indices (Seppälä et al., 2007). This indicates the importance of including particle precipitation in models in order to obtain quantitative agreement with measurements in the mesosphere.

7 Summary

In this paper, we show the first experimental global distribution and seasonal variations of the tertiary ozone maximum in 2002-2006, as obtained from GOMOS data. The tertiary ozone maximum is observed at high latitudes in winter. Peak ozone mixing ratio is from 2 to 4 ppmv, and it can be observed at altitudes 65-75km, depending on latitude and time. The

obtained distributions can be used for validation of global circulation and chemical transport models.

The WACCM model predicts well the position of tertiary ozone maximum and its latitudinal extent, thus confirming general understanding of the processes related to the formation of tertiary ozone maximum. Peak concentrations are slightly overestimated in WACCM compared to GOMOS data. Differences in the model and experimental meteorological fields, inaccuracies in the simulated water distribution in the winter mesosphere or problems with the photolysis rate of water at near-polar night conditions might contribute to the observed discrepancy.

Energetic particle precipitation significantly affects the ozone concentrations at the location of the tertiary maximum (and in the mesosphere in general). Even moderate SPE can destroy the tertiary ozone maximum. TOM mixing ratios are higher during periods of low energetic particle (electrons and protons) precipitation, as was observed in the Northern Hemisphere in Jan-Feb 2006. This indicates the importance of including energetic particle precipitation in models of the middle atmosphere.

Acknowledgements

The work of VS, PV and AS was supported by the Academy of Finland (postdoctoral research projects of VS and AS, THERMES project). The National Center for Atmospheric Research is operated by the University Corporation for Atmospheric Research under sponsorship of the National Science Foundation.

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Figures

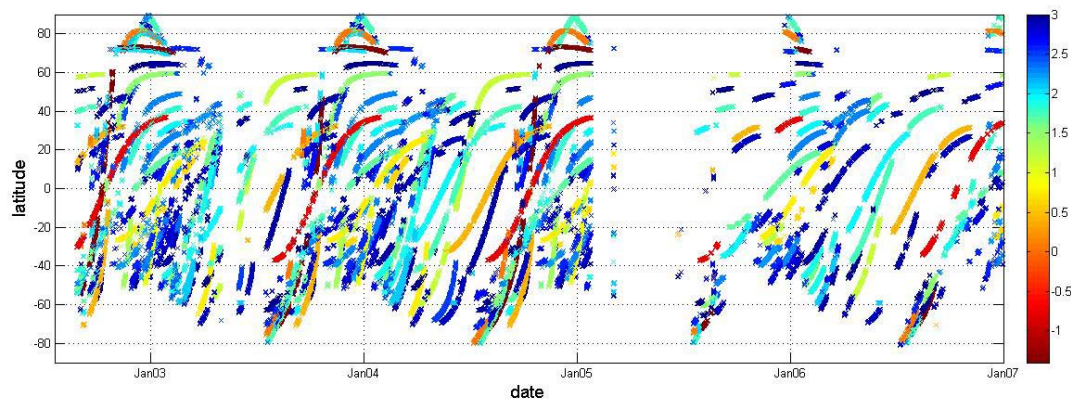


Figure 1 Location of GOMOS nighttime occultations of stars with the effective temperature $\geq 6000\text{K}$ used in the analysis. Color indicates the stellar magnitude: the smaller magnitude, the brighter star.

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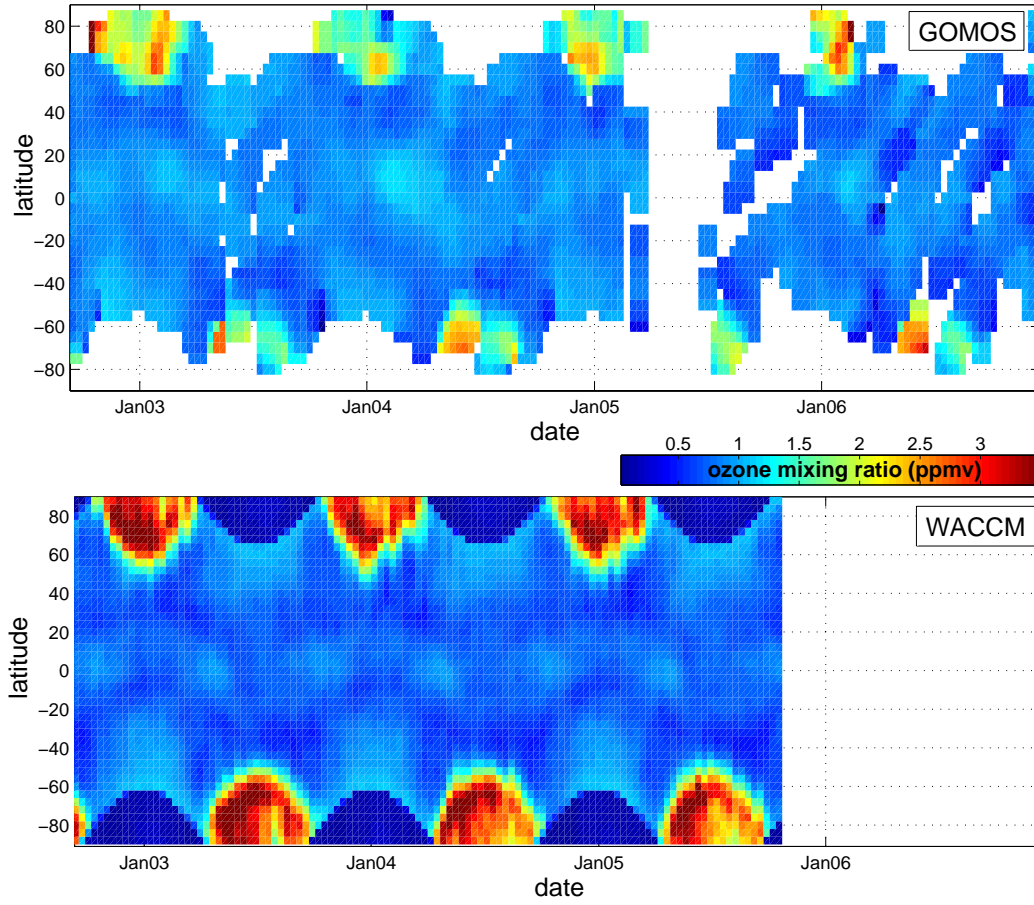


Figure 2 Ozone mixing ratio at 72 km. Top: from GOMOS measurements (zonal mean), bottom: from WACCM simulations (mean over three ensemble simulations for 0° longitude, midnight). Temporal and latitudinal averaging is nearly the same for the both experimental and model data.

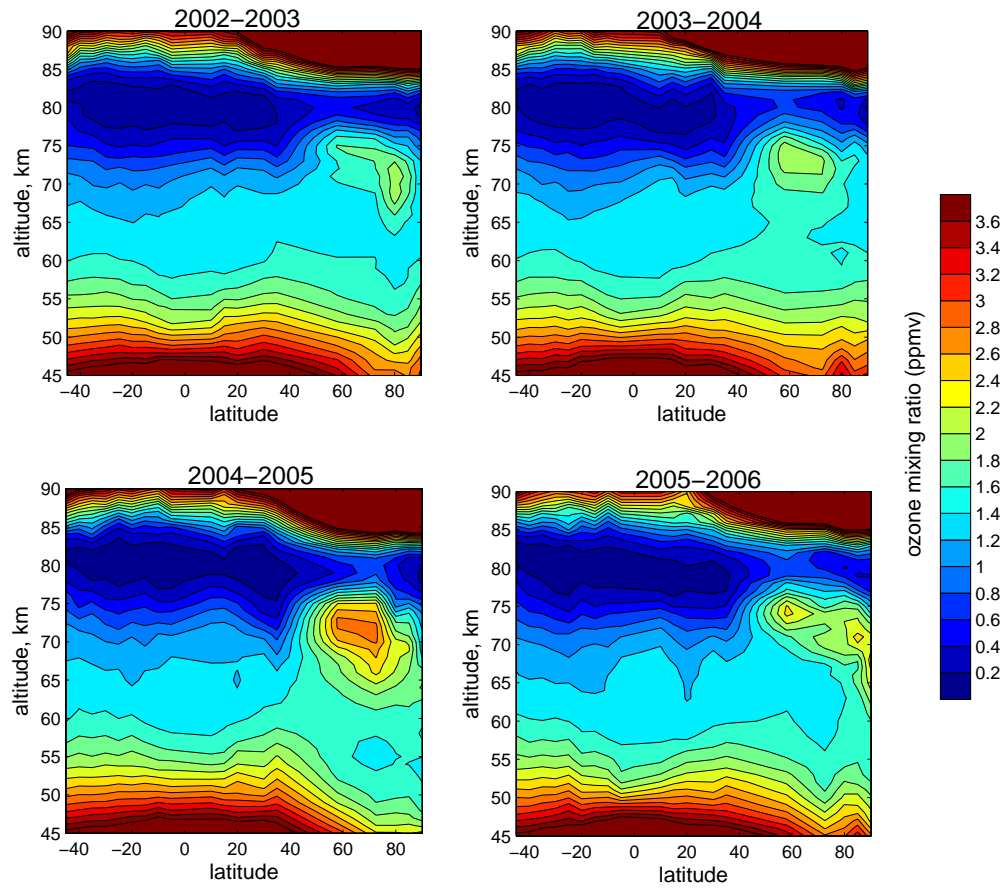


Figure 3 Latitude-altitude section of GOMOS ozone mixing ratio distribution (ppmv) in the longitudinal band from 30W to 30 E averaged for the period 15 Dec- 15 Jan.

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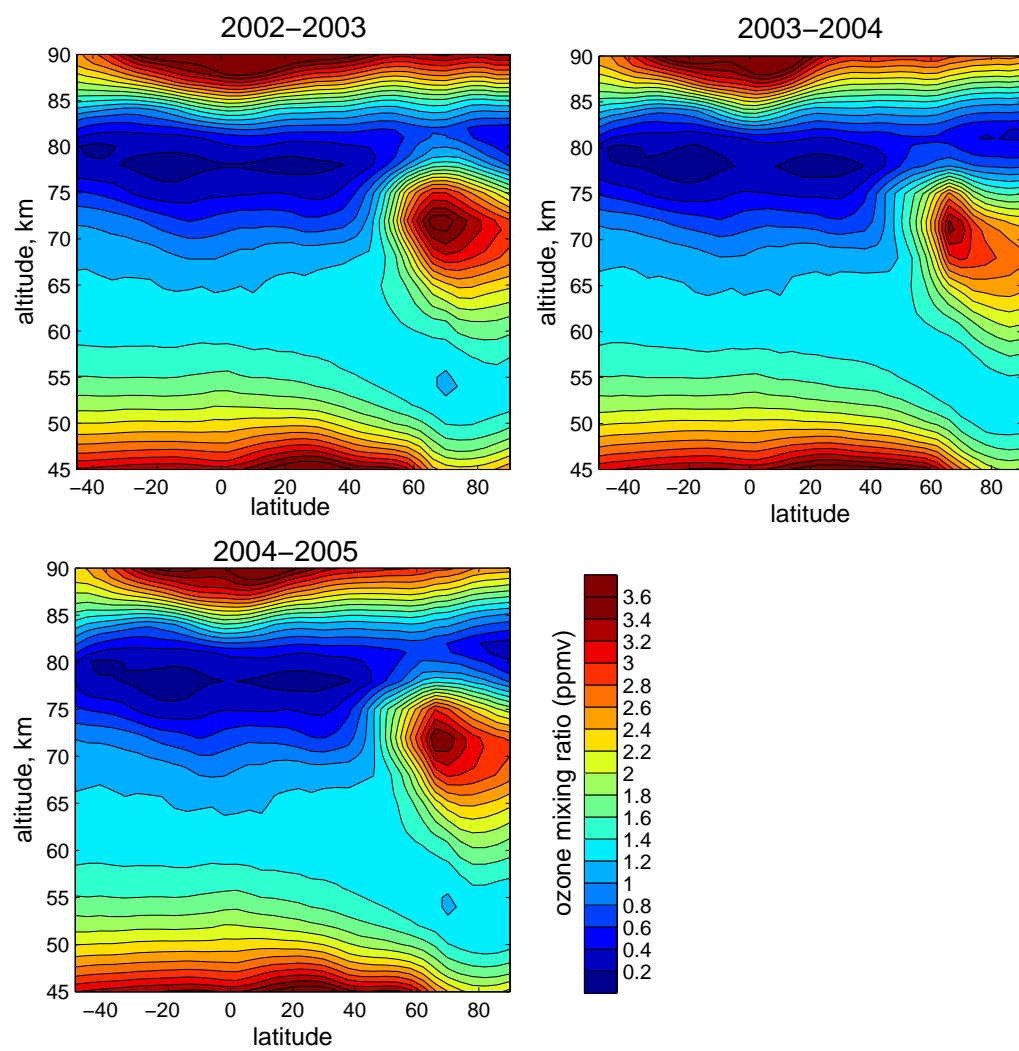


Figure 4 Altitude-latitude sections from WACCM simulations. The data for midnight at longitude 0° are averaged over the period from 15 December to 15 January and over three ensemble simulations.

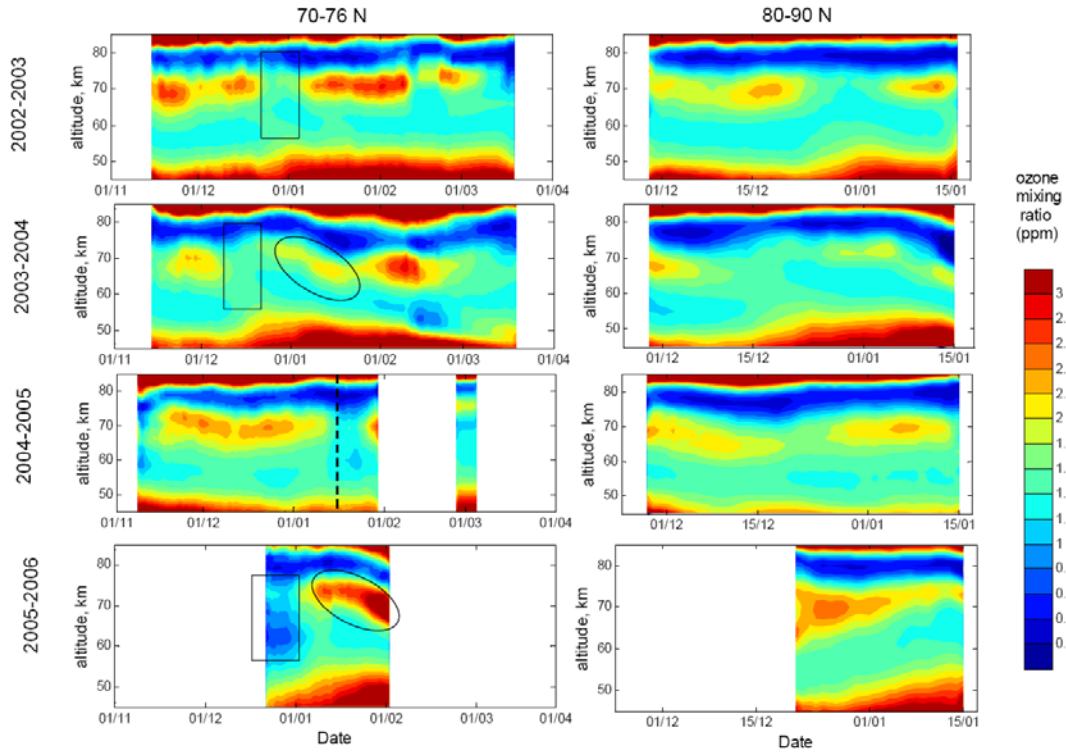


Figure 5 Time series of zonal mean ozone mixing ratio profiles are shown for two latitudinal bands: 70-76 N (left) and 80-90 N (right). 5-day average data are presented. Rectangles indicate periods of sudden stratospheric warmings. Ovals indicate strong downward transport. The dashed line in the subplot for the winter 2004-2005 indicates the onset of SPE in January 2005.

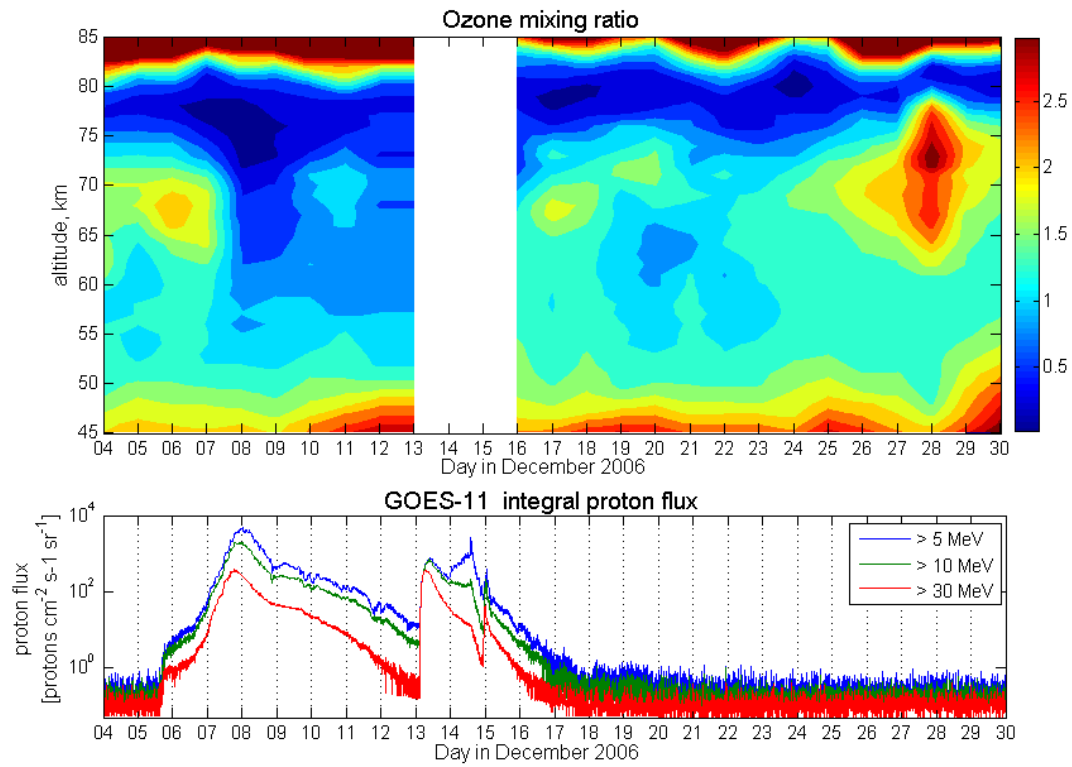


Figure 6 Top: daily ozone mixing ratio profiles at latitudes 60-90N in December 2006. Bottom: proton flux measured by GOES-11. Note the destruction of tertiary ozone maximum caused by the SPE. The Envisat data are unavailable for 13-15 December 2006, due to an anomaly that was probably the result of high geomagnetic activity.