

**January 2005 SPE on  
chlorine species**

A. Damiani et al.

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# Impact of January 2005 solar proton events on chlorine species

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## Abstract

Sudden changes in stratospheric chlorine species in the polar northern atmosphere, caused by the Solar Proton Events (SPEs) of 17 and 20 January 2005, have been investigated and compared with version 4 of the Whole Atmosphere Community Climate Model (WACCM4). We used Aura Microwave Limb Sounder (MLS) measurements to monitor the variability of ClO, HCl, HOCl and Michelson Interferometer for Passive Atmospheric Sounder (MIPAS) on ENVISAT to retrieve ClONO<sub>2</sub>. SPE-induced chlorine activation has been identified. HCl decrease occurred at nearly all the investigated altitudes with the lowest values (of less than 0.25 ppbv) on 21 January. HOCl was found to be the main active chlorine species under nighttime conditions (with increases of more than 0.2 ppbv) whereas both HOCl and ClO enhancements (about 0.1 ppbv) have been observed at the polar night terminator. Further, small ClO decreases (of less than 0.1 ppbv) and ClONO<sub>2</sub> enhancements (about 0.2 ppbv) have been observed at higher latitudes (i.e., at nighttime) roughly above 2 hPa.

While WACCM4 reproduces most of the SPE-induced variability in the chlorine species fairly well, in some particular regions discrepancies between the modeled and measured temporal evolution of the abundances of chlorine species were found. HOCl changes are modelled very well with respect to both magnitude and geographic distribution. ClO decreases are reproduced at high latitudes, whereas ClO enhancements in the terminator region are underestimated and attributed to background variations. WACCM4 also reproduces the HCl depletion in the mesosphere but it does not show the observed decrease below about 2 hPa. Finally, WACCM4 simulations indicate that the observed ClONO<sub>2</sub> increase is dominated by background variability, although SPE-induced production might contribute by 0.1 ppbv.

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## 1 Introduction

Since the early seventies it is known that energetic solar protons are able to affect the neutral chemical components of the polar atmosphere, such as  $[\text{NO}_x] = [\text{NO}] + [\text{NO}_2]$  and  $[\text{HO}_x] = [\text{H}] + [\text{OH}] + [\text{HO}_2]$  via ion chemical reactions (e.g., Swider and Keneshea, 1973; Crutzen et al., 1975; Solomon et al., 1981). These phenomena are important because the  $\text{NO}_x$  and  $\text{HO}_x$  enhancements lead to efficient catalytic cycles of ozone destruction. The short chemical lifetime of mesospheric hydroxyl (OH) radicals causes a relatively brief atmospheric impact from the increased  $\text{HO}_x$  during Solar Proton Events (SPEs) (e.g., Damiani et al., 2010a). A short term ozone depletion occurs from the enhanced SPE-produced  $\text{HO}_x$ , followed by a fast recovery. On the other hand, the longer chemical lifetime of  $\text{NO}_x$ , in particular during the polar winters (when photolysis does not occur), can cause noticeable ozone depletion in the upper stratosphere lasting up to several months (e.g., Randall et al., 2005; Jackman et al., 2009). For this reason, previous efforts have been mainly focussed on this important relationship (i.e.,  $\text{NO}_x$  vs.  $\text{O}_3$ ). Recently, thanks to new satellite observations, many studies have dealt with the impact of SPEs, in particular the so-called Halloween SPE of October–November 2003, on other chemical components (e.g.,  $\text{HNO}_3$ ,  $\text{N}_2\text{O}_5$ ,  $\text{ClONO}_2$ ) (e.g., López-Puertas et al., 2005a,b; Jackman et al., 2005; Rohen et al., 2005; Semeniuk et al., 2005; Seppälä et al., 2006; Egorova et al., 2011; Funke et al., 2011, and references therein). The variability of chlorine species is a further interesting consequence of solar protons entering the Earth's atmosphere. However, this has been investigated only in a few studies (e.g., Solomon and Crutzen, 1981; von Clarmann et al., 2005; Winkler et al., 2009).

Hydrogen chloride (HCl) and chlorine monoxide (ClO) are the main contributors to the  $\text{Cl}_y$  family ( $[\text{Cl}_y] = [\text{Cl}] + [\text{ClO}] + [\text{HCl}] + [\text{ClONO}_2] + [2\text{Cl}_2\text{O}_2] + [\text{HOCl}] + [\text{OClO}] + [\text{BrCl}] + [2\text{Cl}_2]$ ). HCl is the most important stratospheric reservoir species for Cl. ClO carries most of the reactive chlorine in the stratosphere during daytime. Their monitoring is very important since both gases can be considered a significant indicator of the

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effectiveness of the Montreal protocol agreement (e.g., Froidevaux et al., 2006; Jones et al., 2011). On the other hand, a significant interplay between  $\text{NO}_x$  and  $\text{ClO}_x$  is provided by  $\text{ClONO}_2$ , which is another important reservoir species. Also hypochlorous acid (HOCl) has remarkable importance for its catalytic cycle of mid-latitude lower stratospheric ozone loss (Lary, 1997), and because it acts as a link between chlorine and  $\text{HO}_x$  chemistry. However, very few HOCl observations from satellites have been available until recently.

In the upper stratosphere, the chlorine catalytic cycles of ozone destruction have the greatest efficiency around 40 km (Lary, 1997), where SPE-induced changes of chlorine species are expected. There, they can contribute to the short-term  $\text{O}_3$  depletion occurring after a SPE (von Clarmann et al., 2005), which is principally caused by enhanced  $\text{NO}_x$ . In this sense, Salmi et al. (2011) have found that, although not connected to SPEs, the ozone decrease of late February 2009 occurring at polar northern latitudes between about 30 and 50 km, was due to activation of halogen chemistry. Although SPE-induced changes of chlorine species are expected to have an impact only on the short-timescale SPE-induced  $\text{O}_3$  variability, the study of these changes is very useful for validating the chemical schemes used in current atmospheric models.

The pioneering work of Solomon and Crutzen (1981) proposed several chemical mechanisms to explain the impact of the SPE of August 1972 on polar atmospheric composition. Two reactions included in those schemes relevant to this study are:



and



The first experimental confirmation of Reaction (R1), i.e. of the enhancement of  $\text{ClONO}_2$  under SPE conditions, was reported by López-Puertas et al. (2005a) and von Clarmann et al. (2005). The former examined the  $\text{ClONO}_2$  in the context of  $\text{NO}_y$  components (i.e.,  $\text{HNO}_3$ ,  $\text{N}_2\text{O}_5$ ,  $\text{ClONO}_2$  and  $\text{NO}_x$ ), whereas the latter in the context

of chlorine species (i.e., ClO, HOCl, ClONO<sub>2</sub>). In both works, data from Michelson Interferometer for Passive Atmospheric Sounder (MIPAS), recorded during the 2003 Halloween storm, were used. In particular, von Clarmann et al. (2005) showed an enhancement of ClO and HOCl immediately after the SPE, which was interpreted as being due to the fast reaction:



(with Cl arising from Reaction R2), and a subsequent HOCl formation by:



However, recent modelling results (Jackman et al., 2008), compared with MIPAS data recorded during the Halloween storm, revealed good agreement with HOCl data but poor results for ClO. There is still a further important pathway which can significantly reduce the ClO/HCl ratio in the upper stratosphere and, hence, the sensitivity of ozone to chlorine:



Indeed, although the major products of the OH + ClO reaction are the active HO<sub>2</sub> and Cl radicals, even a small branching ratio may lead to less ozone depletion by chlorine species since Reaction (R5) converts an active form (ClO) into a stable reservoir species (HCl) (e.g., Lipson et al., 1997).

Some preliminary confirmation of the relevant role of Reaction (R2) has been shown by Damiani et al. (2009) using Microwave Limb Sounder (MLS) observations recorded during the SPEs of January 2005. They show an HOCl stratospheric enhancement accompanied by a similar decrease in HCl that clearly demonstrates SPE-induced chlorine activation. The conversion of mesospheric HCl into active chlorine was extensively examined by Winkler et al. (2009) using HALOE data recorded during the SPE of July 2000. Their simulations showed that the magnitude of the HCl loss cannot be explained solely by the gas phase Reaction (R2) and model runs with parameterized production

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of  $\text{NO}_x$  and  $\text{HO}_x$  underestimate the HCl loss. Adding to the model reactions of HCl with negative ions improves the agreement between simulation results and observed data in the mesosphere (above about 55 km) but not in the stratosphere (see also Winkler et al., 2011).

Changes in the concentrations of  $\text{O}_3$ ,  $\text{HO}_x$ ,  $\text{NO}_x$ ,  $\text{HNO}_3$  and in temperature caused by the SPEs of 17 and 20 January 2005 have been reported in recent studies (e.g., Veronen et al., 2006, 2007, 2011a; Klekociuk et al., 2007; Damiani et al., 2008; Seppälä et al., 2008; Orsolini et al., 2009; Becker and von Savigny, 2010; Jackman et al., 2011). In the present paper we study the changes that occurred in a larger set of chlorine species (i.e., HCl, ClO, HOCl and ClONO<sub>2</sub>) during the January 2005 SPEs in the Northern Hemisphere (NH) polar region by using MLS and MIPAS data and Whole Atmosphere Community Climate Model (WACCM) version 4 simulations. In Sect. 2 we describe the MLS-Aura and MIPAS-ENVISAT satellite data used. Section 3 reports an overview of the changes in chlorine species induced by the SPEs of January 2005 and compares them with the year-to-year variability and with the variations caused by the SPEs of December 2006. In Sect. 4 we describe the main characteristics of the version 4 of WACCM adapted to reproduce the impact of SPEs on the chlorine species. The space weather and the meteorological conditions of January 2005 are discussed in Sect. 5. In Sect. 6 the spatial distributions and temporal development of measured trace gas concentrations are compared to model calculations and in Sect. 7 these results are discussed in the context of previous work. In Sect. 8 we conclude which features in the data are well understood, which need further investigation, and how further studies could improve our current knowledge on SPE induced chlorine chemistry.

## 2 Satellite data

The main features of the satellite instruments and respective datasets used in the present study will be discussed in this section.

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## 2.1 MLS

The NASA EOS (Earth Observing System) MLS (Waters et al., 2006) is one of the four instruments on Aura satellite, launched on 15 July 2004 to a sun-synchronous near polar orbit with a 13:45 LT (local time) ascending equator-crossing time. MLS scans the limb of the Earth in the flight forward direction, leading to data coverage from 82° S to 82° N latitude on every orbit, and measures microwave emission in different spectral regions. Vertical profiles are measured every 165 km along the suborbital track. Their horizontal resolution is 200–300 km along-track and 3–9 km across-track. The vertical resolution depends on the species. It is typically 3–4 km in troposphere and stratosphere and it degrades up to 8 km in the mesosphere. MLS records profiles of gas phase chemical abundances, temperature, and cloud ice density. Here we will focus mainly on MLS Version 3.3 ClO, HCl, HOCl, N<sub>2</sub>O and temperature Level 2 Data (available at: <http://mirador.gsfc.nasa.gov/index.shtml>). Detailed validation results for the Version 2.2 MLS measurements are shown in Livesey et al. (2007). Additional information on the differences between Version 2.2 and 3.3 data is reported in Livesey et al. (2011) and in the individual validation papers referenced therein.

MLS ClO data (Santee et al., 2008) have a precision on individual profiles between 0.1–0.3 ppbv and fairly small (lower than 15 %) systematic uncertainties (or good accuracy) in the middle/upper stratosphere. Vertical resolution, determined from the full width at half maximum (FWHM) of the averaging kernel, is about 3–4.5 km. The useful range for a single profile should be between 100 and 1 hPa. MLS HCl data (Froidevaux et al., 2008) have a precision of 0.2–0.7 ppbv in the stratosphere and lower mesosphere and good accuracy (10 %). The vertical resolution is about 3 km in the stratosphere and 4–5 km in the lower mesosphere. The useful range is between 100 and 0.32 hPa. MLS HOCl data (see Livesey et al., 2011) are quite noisy, with single-profile precision of 0.3–0.4 ppbv, and generally require some averaging (e.g., weekly zonal mean) to obtain useful precision. In general, MLS HOCl volume mixing ratios (VMRs) seem to be consistent with MIPAS HOCl in the 10 to 2 hPa recommended MLS vertical range. The

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vertical resolution is about 6 km. The single-profile precision of MLS N<sub>2</sub>O data (Lambert et al., 2007) is about 13–25 ppbv (7–38 %) and the accuracy is about 3–70 ppbv (9–25 %) for the pressure range 100–4.6 hPa. The vertical resolution determined from FWHM of the averaging kernel is 4–6 km. The scientifically useful range of the N<sub>2</sub>O data is between 100 to 0.46 hPa. Since the magnitude of SPE-induced variability is expected to be of the order of (or less than) the individual profile precision, we will focus on averaged daily data.

## 2.2 MIPAS

MIPAS is a mid-infrared Fourier transform limb emission spectrometer designed and operated for measurement of atmospheric trace species from space (Fischer et al., 2008). It is part of the instrumentation of the European Environmental Satellite (ENVISAT) which was launched into its sun-synchronous polar orbit of 98.55° N inclination at about 800 km altitude on 1 March 2002. MIPAS passes the equator in a southerly direction at 10:00 a.m. LT 14 to 15 times a day, observing the atmosphere during day and night with global coverage from pole to pole. The instrument's field of view is 30 km in the horizontal direction and approximately 3 km in the vertical direction. MIPAS operated during January 2005 at a reduced spectral resolution of 0.0625 cm<sup>-1</sup> (unapodized) in terms of full width at half maximum. During this period, MIPAS measurements were performed on 10–13 January (middle atmosphere observation mode), 16–18 January (upper troposphere-lower stratosphere (UTLS) observation mode) and 27–28 January (nominal observation mode). Upper atmospheric observation with a scan range of 40–170 km in terms of tangent heights, taken on 21–24 and 29–30 January, are not considered in this study since they provide only marginal information on ClONO<sub>2</sub>, the MIPAS key constituent of this study. The middle atmospheric ClONO<sub>2</sub> observations taken on 10–13 January (i.e., a week before the SPE event) turned out to be of limited use in providing unperturbed background ClONO<sub>2</sub> distributions for this study because of the high temporal variability of this constituent in the vertical range of interest. In the UTLS observation mode, MIPAS records a rear-viewing limb sequence of 19 spectra,

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corresponding to an along track sampling of approximately 290 km whereas, in the nominal observation mode, it records 27 spectra, corresponding to a sampling of about 410 km. Tangent heights covered the altitude range from 50 down to 6 km in the UTLS mode, and 72 to 7 km in the nominal mode. Trace gas profiles have been retrieved from calibrated geolocated limb emission spectra with the scientific MIPAS level 2 processor developed and operated by the Institute of Meteorology and Climate Research (IMK) in Karlsruhe together with the Instituto de Astrofísica de Andalucía (IAA) in Granada. The general retrieval strategy, which is a constrained multi-parameter non-linear least squares fitting of measured and modeled spectra, is described in detail in von Clarmann et al. (2003).

The ClONO<sub>2</sub> data versions used here are V40\_CLONO2\_100 (UTLS mode) and V40\_CLONO2\_201 (nominal mode), all based on the retrieval setup for the original high spectral resolution MIPAS ClONO<sub>2</sub> measurements described in Höpfner et al. (2007) (see von Clarmann et al. (2009) for the description of the reduced spectral resolution ClONO<sub>2</sub> MIPAS retrievals used in this study). Differences between the two versions are related to the vertical scan range and sampling, implemented in the UTLS and nominal observation modes, and do not affect noticeably the data characteristics in the altitude region of interest (López-Puertas et al., 2005a; Funke et al., 2011). The single measurement precision ranges from 0.06 to 0.12 ppbv, increasing with altitude. Vertical resolution is 5–8 km below 2 hPa and 12–14 km above. Meaningful data is obtained below approximately 0.5 hPa (~ 52 km). Unfortunately no MIPAS HOCl data are available for the period under investigation, because since 2005 MIPAS measures at reduced spectral resolution not sufficient for retrieval of HOCl.

### 3 Changes in chlorine species induced by SPEs and year-to-year variability

The descending phase of Solar Cycle 23 was characterized by the occurrence of some intense SPEs (e.g., SPEs of January 2005, May 2005, September 2005 and December 2006). In particular, the SPEs of January 2005 and December 2006 occurred under

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similar seasonal conditions and impacted mainly the atmosphere of the winter hemisphere at Northern Polar regions. In order to evaluate the importance of the chlorine changes induced by the SPEs of January 2005 compared with the year-to-year variability and with the effects caused by the SPEs of 7 and 13 December 2006, Fig. 1 shows the time series of (top to bottom) daily MLS HCl, HOCl and ClO measurements for the winter months (from December through February) of the years 2005–2011 averaged over 70–82° N at 2 hPa (left column) and 1.5 hPa (right column). The occurrence of the SPEs of January 2005 and December 2006 is shown by the vertical black dashed lines. The time series of MLS HCl at 2 hPa are illustrated in the upper left panel. Note the strong decrease, likely induced by Reaction (R2) and lasting about 10 days, that occurred after the SPEs of 17 and 20 January 2005. The HCl decrease is about 0.3 ppbv, and the minimum values occurred the day after the SPE of 20 January. They are among the lowest HCl values ever recorded by MLS at 2 hPa in NH winter during 2004–2011 years. At lower pressures (up to about 0.1 hPa) these features are similar, with a significant depletion appearing during SPEs and persisting only for a few days. This is not shown in Fig. 1. The first proton injection of 17 January seems to cause an initial HCl decrease, over which an additional depletion caused by the SPE of 20 January is imposed (see Sect. 5.1 for the space weather overview of January 2005 SPEs).

The middle left panel of Fig. 1 shows time series of MLS HOCl measurements as in the upper panel. Due to the low HOCl VMRs in the stratosphere, the sudden SPE-induced enhancement of HOCl due to Reaction (R4) can be easily identified. The HOCl increase (more than 0.2 ppbv) in January 2005 is similar in magnitude to the HCl decrease and suggests SPE-induced chlorine activation. These changes are about 3 times higher than the variability encountered in the other years. High HOCl VMRs are also evident during the SPEs of December 2006 (cyan line) but the corresponding HCl decrease, depicted in the upper panel, is not as evident as in January 2005. However, due to the different proton energy spectra of the SPEs of December 2006 compared with the January 2005 ones, the strongest evidence of the SPE-induced variations is

located at slightly higher altitudes. Indeed, the upper and middle right panels show similar changes of HCl and HOCl for December 2006 and January 2005 at 1.5 hPa. These important changes, highlighted also in December 2006, are additional evidence of the presence of chlorine activation under SPE occurrence.

5 The SPE-induced changes of HCl were more than 0.2 ppbv at 1.5 hPa, but even larger HCl decrease took place in February 2006, 2009 and 2010 corresponding to periods with strong winter polar vortices in the upper stratosphere following intense stratospheric sudden warming (SSW) events that occurred in January of those years (e.g., Siskind et al., 2007; Manney et al., 2009; Damiani et al., 2010b). Further HCl  
10 changes (enhancements) occurred during some SSW events (see for example the HCl peak at 2 hPa in late February 2008). Figure 1 shows that HCl depletions were coupled with intense HOCl enhancements only during SPEs whereas negligible HOCl variations occurred after or during SSWs when the active chlorine was mainly in the form of ClO. On the other hand, no clear evidence of ClO changes connected to the  
15 SPEs of January 2005 and December 2006 is present at 2 hPa. Limited ClO decreases are present at 1.5 hPa even if they cannot be attributable with certainty to SPE occurrence. In addition, a sudden spike of ClO VMR (about 0.1 ppbv) is visible at 1.5 hPa on 19 January 2005. We will examine these issues in Sect. 6.

#### 4 Description of the whole atmosphere community climate model (WACCM4)

20 In the following, MLS and MIPAS data are compared with version 4 of the National Center for Atmospheric Research (NCAR) Whole Atmosphere Community Climate Model (WACCM4) results ([http://www.cesm.ucar.edu/working\\_groups/WACCM/](http://www.cesm.ucar.edu/working_groups/WACCM/)). An earlier version of the model, WACCM3, was used extensively to investigate the atmospheric effects of SPEs (Jackman et al., 2008, 2009, 2011). The present model  
25 version, WACCM4, was used in the recent Funke et al. (2011) study of the 2003 Halloween SPEs. For this investigation, WACCM4 is forced with meteorological fields from the Modern Era Retrospective-analysis for Research and Applications (MERRA,

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<http://gmao.gsfc.nasa.gov/research/merra/>). The same approach was used in the study of the 2003 Halloween SPE (Funke et al., 2011) and the study of the 2006 stratospheric sudden warming (Marsh, 2011). As in those studies, the forcing is achieved by relaxing horizontal winds and temperatures with a time constant of approximately 50 h from the surface to 40 km. Above that level the forcing is reduced linearly, so that the model is free-running between 50 km and the model top. The model has 88 levels in the vertical, with a resolution in the horizontal of 2.5 longitude by 1.9 latitude.

Two sets of simulations were conducted: one with hourly energetic proton ionization rates derived from GOES 11 flux data, and a reference run without ionization. Ionization rates are shown in Fig. 2 and discussed in Sect. 5.1. Neither simulation included highly energetic electron precipitation. WACCM4 output closest to each of the satellite (MLS and MIPAS) geo-locations in space and time has been used. It should be noted that in order to obtain more realistic comparisons, MLS or MIPAS averaging kernels (AK) have been applied to WACCM4 output.

## 5 Space weather and meteorological conditions during January 2005

Some features of the space weather and the meteorological conditions of the upper stratosphere/lower mesosphere at Northern high latitudes during January 2005 are discussed in the present section.

### 5.1 Space weather

January 2005 was characterised by low/moderate solar activity between 1 and 13 January and after 23 January, while during 15–20 January, the active region 10720 produced five powerful solar flares classified as X1.2, X2.6, X3.8, X1.3 and even X7.1 (see for example The Preliminary Report and Forecast of Solar Geophysical Data, called “The Weekly”, available at <http://www.swpc.noaa.gov/weekly/index.html>).

Three SPEs occurred on 16, 17 and 20 January. In the following, we will refer exclusively to the SPEs of 17 and 20 January. The 17 January event was

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characterised by a large injection of solar particles peaking at 5040 pfu (Particle Flux Units;  $1 \text{ pfu} = 1 \text{ p cm}^{-2} \text{ sr}^{-1} \text{ s}^{-1}$ ) at the 10 MeV channel of the National Oceanic and Atmospheric Administration (NOAA) Geostationary Operational Environmental Satellite (GOES), and 20 January exhibited the hardest and most energetic particle event of the solar activity cycle 23 (see for instance Damiani et al., 2008). Indeed, on 20 January the peak for the 100 MeV channel of GOES was 652 pfu (we recall that during the Halloween Storm the peak flux at the 100 MeV channel was 155 pfu), while the 10 MeV peak was 1860 pfu (a list of SPEs affecting the Earth's environment is provided by NOAA at <http://www.swpc.noaa.gov/ftplib/indices/SPE.txt>). The different energy spectra of the two events lead to impacts at quite different layers of the polar atmosphere. In particular, the SPE of 20 January is expected to influence the atmosphere to a deeper level than the 17 January event due to its larger flux of protons at high energy.

It is important to state that the ion production rate was mostly smaller in both 17 and 20 January 2005 compared with the SPEs of October 2003 and July 2000. For example, the daily average peak of ion pair production rates was more than  $5000 \text{ cm}^{-3} \text{ s}^{-1}$  in October 2003 and July 2000, whereas it was about  $1000 \text{ cm}^{-3} \text{ s}^{-1}$  on 17 January 2005 (Jackman et al., 2008, 2011). The largest ion production rate in 2005 was located at higher altitudes compared with 2003. That influenced the mesospheric  $\text{NO}_x$  production by moving the SPE-induced peak of  $\text{NO}_x$  in January 2005 toward higher altitudes compared with 2003. Moreover, since the SPE of 20 January was characterised by a very short proton injection, it is important to take into account a finer temporal resolution of the ion production rate. Therefore, Fig. 2 shows the hourly average ionization rate, from 100 hPa to 0.001 hPa, as a function of time for 16–22 January 2005. The proton flux data from GOES-11 were used to compute ion pair production profiles used in the model (see Jackman et al., 1980, for the methodology). We can note some structure that is not resolvable when looking at the daily average ionization rate (compare this figure with Fig. 1 of Jackman et al., 2011). The most intense ionization, with peak greater than  $2000 \text{ cm}^{-3} \text{ s}^{-1}$ , is above 1 hPa on 17 January whereas it is below 1 hPa on 20 January. However high ion rates last only a few hours on 20 January.

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## 5.2 Meteorological conditions

The absence of strong temperature increases and associated reversals of the zonal winds indicates that the NH winter of 2004/05 was without major sudden stratospheric warmings (e.g., Chshyolkova et al., 2007; Xu et al., 2009; Damiani et al., 2010b). However, the winter of 2004/05 was particularly cold and dynamically active, especially in the lower stratosphere (Manney et al., 2006; WMO, 2007; El Amraoui et al., 2008). The temperatures exhibit rapid fluctuations in some periods, for example at the beginning and at the end of January. While the polar vortex reached its strongest state and was centred on the pole in December, it was elongated, slightly shifted off the pole and had a decreased extension in the lower stratosphere during these disturbed periods of January (Chshyolkova et al., 2007). Because one of these periods occurred a few days after the SPEs occurrence, this anomalous dynamical vortex behavior likely weakened the chemical impact of the SPEs.

Since meridional transport and mixing can play a key role in the magnitude and spatial distribution of SPE-induced variations, Fig. 3 reports the zonal mean geographic distribution of  $N_2O$  in four periods (from top to bottom: 12–16 January, 17–21 January, 22–26 January and 27–30 January) during the second half of January 2005 as recorded by MLS (left panels) and simulated by WACCM4 (right panels). Due to the strong variation of  $N_2O$  with altitude, we show the  $N_2O$  relative anomaly, with respect to the meridional average (40–82° N), expressed in percent. Lower abundances of  $N_2O$ , surrounded by a pronounced latitudinal gradient, indicate the position of the polar vortex. During the period of the SPEs (17–21 January) MLS shows that the vortex is well defined in the stratosphere up to 2–3 hPa, whereas above this layer, a smaller latitudinal gradient is indicative of more intense latitudinal redistribution. This is roughly in agreement with the analysis of the NH polar vortex dynamics of Verronen et al. (2011a) showing increasing horizontal transport and mixing over the altitude range from the lower ( $\sim 1$  hPa) to middle mesosphere during the January 2005 SPEs. Note that the area 70–82° N, to which we will refer in the paper later on (see, for example, Fig. 4),

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seems to be well inside the vortex (with the consequent weak redistribution) only up to about 3 hPa. Above this pressure level, the vortex gets weaker particularly during the following days (lower panels). Therefore SPE-induced effects should be longer lasting in the middle stratosphere compared with the effects in the upper stratosphere/lower mesosphere. WACCM4 reproduces the MLS N<sub>2</sub>O pattern quite well but simulates weaker latitudinal redistribution in the mesosphere.

The temporal evolutions of daily mean changes of temperature and N<sub>2</sub>O with respect to January average, averaged over 70–82° N and 58–70° N, are reported in the upper panels of Figs. 4 and 5. The patterns of the MLS temperature and N<sub>2</sub>O changes (left panels) are quite similar to each other and show two major variations at the beginning and the end of January connected to the disturbed polar vortex periods. In general higher temperatures and N<sub>2</sub>O values are associated with entrainment of external vortex air into the 70–82° N region. Thus they can indicate intrusions or vortex displacement. However, the isolines of the absolute N<sub>2</sub>O abundances (not shown) suggest continuing confined air descent in the middle stratosphere throughout January. Therefore in this region the vortex mixing barrier should remain quite strong at this time according to Chshyolkova et al. (2007). More frequent variations of temperature and N<sub>2</sub>O occur in the upper stratosphere, so, as already discussed, this region of the vortex is less stable compared with lower altitudes. Further, in Fig. 4 we observe an increase of N<sub>2</sub>O values in the upper stratosphere starting around 18 January and becoming more pronounced towards the end of the month. Inspection of the MLS N<sub>2</sub>O data maps (not shown) suggests that, although the polar vortex remains substantially well defined in January, it is elongated and shifted away from the North Pole on many days. In particular, after 27 January, a considerable part of the region inside 70° N is occupied by extra-vortex air and obviously that influences the daily means. The WACCM4 simulations follow the changes in MLS temperature and N<sub>2</sub>O quite well, although some discrepancies in their magnitude are visible in the second half of January. Finally, zonal mean changes of temperature and N<sub>2</sub>O on 17–21 January with respect to 12–16 January at 40–90° N are reported in the upper panels of Fig. 6. Temperature and N<sub>2</sub>O patterns are consistent

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with the general behaviour of the vortex already depicted, hence we observe its intensification in the middle stratosphere and its weakening in the upper stratosphere.

## 6 Changes in chlorine species: satellite data vs. WACCM4

Figure 4 shows the temporal evolution of daily mean changes of HCl, HOCl, ClO with respect to January average, averaged over 70–82° N, therefore under almost complete nighttime conditions (see Sect. 5.2 for the description of temperature and N<sub>2</sub>O panels). The loss in MLS HCl (left panel) occurred after 17 January and affects nearly all investigated pressures. The strongest decrease (of less than 0.25 ppbv) occurs uniformly in a rather wide region ranging from 1.5 to 6 hPa on 21–22 January. This decrease lasts until the end of January slightly descending to around 6–8 hPa. WACCM4 simulations of HCl (central panel) show a quite good agreement in the lower mesosphere but not in the stratosphere. The HCl decrease below 2 hPa, well evident in MLS data, is not reproduced by WACCM4 results. The right panels of Fig. 4 (WACCM b-p) show the difference between the perturbed (with SPEs) and the base (without SPEs) WACCM4 simulations. Comparing WACCM4 results with WACCM(p-b) shows that, in WACCM4 results, the HCl decrease caused by the SPEs between 2 and 4–5 hPa is masked by background variability, and the HCl depletion below 4–5 hPa is not ascribable to SPE impact. The MLS HOCl temporal evolution shows an increase of more than 0.2 ppbv on 18–21 January 2005. After the first enhancement caused by the SPE of 17 January (more than 0.15 ppbv), a further HOCl increase is induced by the SPE of 20 January (see Fig. 4) and HOCl disturbances are slightly moved downward because of the harder spectrum of the second SPE. This enhancement is clearly evident above 4–5 hPa and lasts up to about 25 January. Both the pattern and the magnitude of these changes are well reproduced by WACCM4.

The lower left panel of Fig. 4 shows the temporal evolution of the MLS ClO. ClO presents a weak increasing trend in the region 3–5 hPa, that reaches the highest values after 20 January, and a decrease in the upper stratosphere (1–2 hPa) which is strongest around 18 January. These changes seem to be mainly related to dynamical variability.

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The decrease of ClO starts before the occurrence of the SPEs and therefore it masks the influence of the solar forcing. However, a further depletion in ClO occurs on 17 and 18 January. Then, on 19 January a slight ClO rise is quite evident, followed by another depletion on 20 and 21 January. This behaviour is compatible with a reduction of ClO via Reaction (R4) and the consequent HOCl formation. Therefore, active chlorine exists mainly in the form of HOCl above approximately 3–4 hPa at 70–82° N (i.e., at nighttime conditions). The WACCM4 result (central panels) reproduces the general features of MLS ClO. However, the onset of the particle-induced ClO depletion is more clearly visible since the observed decreasing trend before the SPEs is less pronounced in the simulation. The rather constant values of the WACCM(p-b) differences (right panel) below ~ 4 hPa suggest that the observed and modelled ClO increase is not connected to the SPE.

Figure 5 shows the temporal evolutions as in Fig. 4 but with values averaged over 58–70° N, and therefore roughly under twilight conditions. In this geographic region MLS HCl still undergoes a significant decrease (up to about 0.2 ppbv) that is mostly evident above 1 hPa and around 5 hPa. As for the 70–82° N region (Fig. 4), WACCM4 does not reproduce the HCl depletion below about 2 hPa. Moreover, unlike the Fig. 4, it clearly underestimates the mesospheric HCl depletion. From the comparison of WACCM4 results and WACCM(p-b) differences it follows that the simulated HCl decrease below 4 hPa is not due to SPE.

It is interesting to examine the short-term variations of HOCl and ClO as observed by MLS during the SPEs occurrence. At 58–70° N, under greater solar illumination with respect to the 70–82° N region, an HOCl enhancement, with peaks of about 0.1 ppbv on 18, 20 and 22 January, is evident around 2–3 hPa. ClO changes are also evident. In general, a ClO decrease during SPE-influenced days, as seen in Fig. 4, is not present. On the contrary, a blurred ClO enhancement takes place around 1 hPa. Moreover, ClO increases are quite evident on 19 and 21 January, i.e. not in conjunction with but the day after the HOCl peaks (note that the ClO peak of 19 January is evident also at nighttime, see Figs. 1 and 4). This suggests that the SPE-generated HOCl is photolyzed after the

SPEs, leading to ClO formation via Reaction (R3). In addition, somewhat smaller ClO increases seem to be present also around 4–6 hPa, roughly coincident with the region where stratospheric HCl decrease occurred. WACCM4 reproduces very well the HOCl increases whereas it does not show any ClO enhancement. On the contrary it shows a very small ClO decrease (as at nighttime), coupled with the SPEs occurrence, that is not present in the MLS data.

Figure 6 illustrates the zonal mean changes of (from top to bottom) HCl, HOCl, ClO on 17–21 January 2005 with respect to 12–16 January 2005 at 40–90° N (see Sect. 5.2 for the description of temperature and N<sub>2</sub>O panels). The geographic distribution of MLS HCl and HOCl changes (left panels) shows that the perturbations above 2 hPa extend well also into mid-latitudes. This feature is less pronounced in WACCM4 (central panels), most probably related to the underestimation of the latitudinal redistribution at higher altitudes. Also below 4 hPa negative changes of MLS HCl extend towards lower latitudes, however they are not connected to SPEs occurrence. N<sub>2</sub>O changes at 50–70° N (evident in the above panels) are indicative of the different behavior of the polar vortex in these two periods. It was roughly circular in shape and centered over the pole on 17–21 January whereas it was elongated and slightly shifted off the pole on the days before. Therefore, taking into account differences in HCl abundance between inside and outside of the polar vortex, dynamics is likely the main reason of the negative changes in HCl at 50–70° N visible in both observations and modelling results.

Both HCl and HOCl show strongest changes at the highest latitudes, where the photolysis is lowest and the vortex is more isolated. Because of the limited ionization at lower altitudes, only HCl changes above about 4 hPa are likely due to SPEs occurrence (see also the WACCM(p-b) in the right panel of Fig. 6). On the other hand, the evident model-data HCl discrepancies below 4 hPa inside the vortex could hardly be attributed to unrealistic model dynamics, given the good agreement of observed and modeled tracer (and temperature) evolutions, there. However, it could be possible that this inconsistency is caused by a different chemical response to dynamical background variations in the model.

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The zonal mean change of CIO (lower panels of Fig. 6) shows a well defined pattern in both MLS data and WACCM4 results. The observed CIO depletion of about 0.1 ppbv is present only at high latitudes whereas a CIO enhancement of a similar magnitude is visible at about 60–70° N. WACCM4 results underestimate the response at lower latitudes and overestimate the vertical extent of the CIO decrease at high latitudes. The different response of CIO to SPEs as function of latitude seems to be related to the presence/absence of solar illumination. Therefore, HOCl is the main active chlorine species during nighttime, whereas both CIO and HOCl show perturbations, probably due to SPEs, in twilight regions. However, WACCM(p-b) differences suggest that the CIO increase at lower latitudes is not linked to SPE occurrence (see also the WACCM(p-b) of Fig. 5). We will reexamine these features in the discussion section.

MIPAS and WACCM4 CIONO<sub>2</sub> zonal mean changes with respect to 16–17 January 2005 are shown in Fig. 7 for 18 and 27 January 2005. In the upper panels of Fig. 7 we can see that the zonal mean change on the day after the first SPE does not indicate evidence for a significant SPE-induced CIONO<sub>2</sub> response; although small changes on the order of 0.03–0.04 ppbv are observed (left panel). Changes of similar magnitude are also simulated by WACCM4 (central panel). The lower panels show that the zonal mean change on 27 January is characterised by large CIONO<sub>2</sub> changes (~ 0.2 ppbv) around 1 hPa and by a smaller variation (~ 0.07 ppbv) at about 3 hPa. WACCM4 results reproduce well the observations in terms of magnitude and altitude of these peaks. In addition, WACCM(p-b) (right panel) differences suggest that the CIONO<sub>2</sub> changes around 3 hPa are mainly attributable to background variations, whereas the enhancement at 1 hPa is induced by both background variability and SPE occurrence (with the latter contributing by ~ 0.1 ppbv). Also in this case, the more intense latitudinal redistribution in the upper stratosphere is probably the cause of the higher CIONO<sub>2</sub> increases outside the polar regions at these altitudes.

## 7 Discussion

The analysis presented above highlighted the presence of chlorine activation under SPE conditions although some aspects still remain unclear. Therefore, it is worthwhile to discuss these results in the light of past findings. Some important features distinguish the January 2005 SPEs from other events (e.g., SPEs of July 2000 and October–November 2003) investigated in the context of their impact on chlorine. First, the polar night region was much more extended in January 2005 compared with October 2003 and July 2000. The well established vortex coincided with the terminator in 2005, while in 2003, the early winter vortex was partially illuminated, and the SPE of July 2000 was investigated only in the summer hemisphere. It is also important to emphasise the smaller ionization rates modeled for January 2005 compared with other events (see Sect. 5.1), which caused a smaller production of  $\text{NO}_x$  at upper altitudes in 2005 compared to 2003. As we will see in the following, these elements are important in order to explain the impact of the investigated events on the atmospheric composition.

To our knowledge, only Winkler et al. (2009) have carefully examined the variation of HCl under SPE conditions. However, the different time of the year of the examined SPE (July 2000) and the analysis performed by means of solar occultation data (HALOE) and a 2-D chemistry model (with ion chemistry) prevent a full comparison of their results with ours. Some differences in the vertical distribution of the HCl depletion, caused by the different SPEs, are evident. HALOE data in 2000, recorded in the illuminated NH, show variability between 65 km ( $\sim 0.1$  hPa) and 35 km ( $\sim 5$  hPa), with the maximum stratospheric HCl depletion (about 0.4 ppbv) at 50 km and a roughly constant reduction of this enhancement down to 35 km, where the HCl changes are small ( $\sim 0.05$  ppbv). In contrast, MLS data in January 2005 show the maximum HCl loss (of less than 0.25 ppbv) uniformly distributed in a rather wide region ranging from 1.5 to 6 hPa. These differences are mainly related to the stratospheric background variations in January 2005, to the different atmospheric seasonal conditions at Northern Polar

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regions in July and January, and to the occurrence of two successive SPEs in January 2005.

In fact, it is important to identify the distinct contributions arising from different sources to the HCl decrease in January 2005. In January 2005 we have the overlap of the effects of two SPEs, with the second one, characterized by a harder spectrum, able to impact the already low stratospheric HCl values persisting longer at nighttime (Fig. 4). On the contrary, only one SPE (even if more intense) occurred in July 2000. The diverse proton energy spectra during the two SPEs in January 2005 is reflected in the shape of the HCl decrease. The first depletion ( $\sim 0.2$  ppbv), due to 17 January SPE, affected mesospheric altitudes on 18 January. In this case, the location of the maximum HCl depletion (around 1 hPa) is roughly in agreement with previous results. Then, a second and more intense stratospheric decrease, superimposed over the previous one, occurred after the 20 January SPE. This strongest decrease occurs rather uniformly in the upper and middle stratosphere on 21–22 January and it lasts until the end of January around 6–8 hPa. However, because of the lack of significant ionization at these altitudes, this long lasting tail of HCl decrease seems to be not connected to the SPE occurrence. The expected low SPE-induced impact below about 4 hPa is likely superimposed over background variations (see Sect. 6). It results in a stronger than expected HCl decrease with a maximum around 5 hPa on 21–22 January and in the evident HCl tail persisting until the end of January. Hence, these model-data discrepancies, mainly present at about 5–8 hPa (see Figs. 4 and 5), are mostly due to background atmospheric variability. Their investigation is beyond the scope of this study and they will require a specific analysis.

An additional element that could play a role in explaining these differences in the HCl decrease is related to the different background of OH VMRs in January and July in Northern Polar regions. The OH VMRs in the stratosphere and mesosphere depend mainly on the Solar Zenith Angle (SZA). In the upper stratosphere, between 7 hPa and 1 hPa, the typical daytime OH abundance is about 0.1–0.6 ppbv, with lower values at lower altitudes, whereas during the night the OH VMRs are uniformly very close to zero.

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In order to have appreciable HCl changes, the SPE-induced OH increase has to be, at least, of the same order of magnitude (or larger) as the OH background. Therefore differences in the vertical distribution of the HCl depletion, caused by the SPEs of July 2000 and January 2005, had likely been influenced by differences in the background of OH VMRs.

A further difference between the SPEs of January 2005 and July 2000 is the duration of the mesospheric HCl decrease. Indeed, HCl recovered in a couple of days in July 2000, whereas the depletion lasted more than one week in January 2005. Obviously the occurrence of two successive SPEs in January 2005 is the first reason of the correspondent longer HCl depletion. Moreover, the confinement maintained by the polar vortex in January 2005 and the absence of photolysis at high latitudes, that likely prevented a quick shift towards HCl from the SPE-induced active forms (e.g., the destruction of HOCl by photolysis), are additional causes compared to the polar summer conditions in July 2000.

Winkler et al. (2009) and Winkler et al. (2011) showed that simulations using parameterized HO<sub>x</sub> production rates and Reaction (R2) underestimate the HCl depletion in the mesosphere and stratosphere, and the inclusion of ion chemistry significantly improves the results in the mesosphere but not in the stratosphere. WACCM4 reproduces quite well the HCl depletion in the lower mesosphere (about 0.2 ppbv) at nighttime on 18 January. Also at the upper mesospheric altitudes where the HCl depletion is even more intense (not shown), WACCM4 results are in fair agreement with experimental data without the necessity of including ion chemistry. However, WACCM4 underestimates the mesospheric HCl response at 58–70° N (see Fig. 5). This could be related to the underestimation of the redistribution of the bulk of depleted HCl molecules originated mainly at higher latitudes.

In the upper stratosphere (roughly above 2–3 hPa), where HCl changes are related to the SPEs, the model tends to underestimate the observed depletion (see Figs. 4 and 5). The lack of HCl loss (or its underestimation) in the WACCM4 results in this region is in line with the difficulties encountered by Winkler et al. (2011) when modeling

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the stratospheric HCl response during the July 2000 event. The cause is not easily understandable. Possible reasons that could contribute to explain such difference are:

1. SPE-induced ionization included in the model is too small, either due to an underestimation of proton fluxes or due to the neglect of electron-induced ionization (Verronen et al., 2011b). However, MLS OH is reproduced by WACCM4 quite well down to 3 hPa (not shown), so this possibility appears to be unlikely.
2. The insertion of ion chemistry in the model.
3. The uncertainty of the rate constant of Reaction (R5) included in the model with the JPL rate constant (Sander et al., 2006). It is connected to the deactivation of ClO into the reservoir HCl but probably could play only a limited role in this discrepancy.

Further analyses will be necessary in order to explain the model data discrepancies.

Also the behaviour of ClONO<sub>2</sub> deserves special attention. Figure 7 shows that ClONO<sub>2</sub> changes are not present on the day after the first SPE. On 27 January we note an enhancement of ~0.2 ppbv at about 1 hPa. The work of Funke et al. (2011), who use a reprocessed version of the ClONO<sub>2</sub> data set presented by von Clarmann et al. (2005), shows the SPE-induced ClONO<sub>2</sub> peak (~0.4 ppbv) at about 3 hPa in October-November 2003. Either the altitude or the magnitude is somewhat different compared with the peak value observed on 27 January 2005. The ClONO<sub>2</sub> formation depends on the availability of NO<sub>2</sub> and ClO, and on the density. In January 2005 the SPE-induced production of odd nitrogen was considerably lower compared to the October-November 2003 events. Jackman et al. (2008) estimated that the NO<sub>y</sub> production in 2005 was ~4 times smaller than in 2003, which could lead to a lower production of ClONO<sub>2</sub> in January 2005. Moreover, we could expect a ClONO<sub>2</sub> enhancement at somewhat lower altitudes compared with the altitudes of the SPE-induced NO<sub>2</sub> layer (because of the pressure dependence of ClONO<sub>2</sub> formation). Because of the different distribution of the ionization (see Sect. 5.1), the SPE-induced NO<sub>2</sub> formation in January

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2005 should occur mainly at higher altitudes compared with 2003. Hence, it is likely that the  $\text{ClONO}_2$  increase ( $\sim 0.2$  ppbv) of 27 January at about 1 hPa (Fig. 6) is due, at least partly, to the SPEs. This seems compatible with the incomplete recovery of HCl around 1 hPa lasting until the end of January (see Figs. 4 and 5) caused by an initial chlorine activation ( $\text{HCl} \rightarrow \text{ClO}$ ) and a further possible deactivation into the  $\text{ClONO}_2$  reservoir. The elevated altitude of  $\text{HNO}_3$  enhancements that occurred during the SPEs of January 2005 and recently reported by Verronen et al. (2011a) (with peak around 60 km, so at altitudes higher than previous observed  $\text{HNO}_3$  variations caused by SPEs (e.g., López-Puertas et al., 2005a)) is a further element corroborating our finding of SPE-induced  $\text{ClONO}_2$  changes at elevated altitudes.

It is also interesting to examine in detail the behaviour of ClO and HOCl. von Clarmann et al. (2005) showed an increase of MIPAS ClO VMRs immediately after the SPE of September-October 2003. A recent analysis based on reprocessed IMK/IAA MIPAS data (Funke et al., 2011) clarified that MIPAS recorded increased ClO (about 0.1 ppbv) mainly outside the polar night region, whereas a ClO decrease (about 0.2 ppbv) occurred in the polar night region. This is consistent with the MLS results presented in Fig. 6 where the ClO decrease is roughly confined to the polar night sector and increased ClO is evident at the terminator (in particular above 2 hPa at  $60\text{--}70^\circ\text{N}$ ). However, WACCM4 simulations indicate that only the ClO decrease is caused by SPEs. On the other hand, the strong HOCl enhancement is well reproduced by WACCM4 at nighttime conditions.

During the daytime, HOCl is quickly destroyed by photolysis



and by reaction with OH (or atomic oxygen):



Funke et al. (2011) suggested that, because OH is the dominant  $\text{HO}_x$  constituent during daytime (at least in the upper stratosphere; below roughly 5 hPa  $\text{HO}_2$  starts to be

comparable to OH and gets even larger than OH at lower altitudes) and Reactions (R6) and (R7) are the major daytime processes, active chlorine is mainly in the form of ClO under these conditions. The ClO enhancements in MIPAS data outside the polar region during the Halloween storm could then be related to SPE-induced chlorine activation if the production of HO<sub>x</sub> was above the typical HO<sub>x</sub> background levels. In January 2005 MLS data indicate that the SPE-induced OH rise is comparable with background levels (about 0.05 ppbv) at 1 hPa and 65–70° N, coinciding with the altitude of the ClO peak (see left panel of Fig. 6), whereas there are not appreciable changes at lower altitudes.

Figure 6 shows that the region of low HCl values extends towards low latitudes in the upper (~ 1 hPa) and middle stratosphere (~ 7 hPa), whereas it remains more confined inside high latitudes around 3 hPa. The MLS ClO structure in the same region (roughly 60–70° N) presents higher ClO values at the upper altitudes, where there is availability of active chlorine molecules arising from the HCl depletion via SPE-induced OH, whereas little change was recorded around 3 hPa. Further, HOCl enhancements are present also at 60–70° N above about 3 hPa (Fig. 6). Increased production of HOCl in the middle stratosphere, induced by accelerated odd hydrogen chemistry (via Reaction R4), has been reported by von Clarmann et al. (2009) in association with a displacement of the polar vortex towards lower latitudes during a SSW event. SPE occurrence represents a further example of accelerated odd hydrogen chemistry as indicated by both HOCl and ClO increases between 58–70° N (see Fig. 5). Comparing the temporal evolution of ClO and HOCl zonal means, it follows that ClO peaks seem to temporally follow HOCl changes. This suggests that SPE-induced HOCl increases may have been photolyzed leading to the shift from HOCl to ClO. The dynamically active polar vortex could also have contributed to complicate the situation by masking and mixing chlorine changes. However, WACCM(p-b) of Figs. 5 and 6 does not show any SPE-induced ClO enhancements whereas it reproduces well HOCl changes also outside the polar night terminator.

Considerable uncertainty persists in laboratory measurements of the rate constant for the Reaction (R4). Some studies (e.g., Kovalenko et al., 2007; von Clar-

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mann et al., 2009) pointed out the necessity of increasing it towards the value from Stimpfle et al. (1979) that is about a factor-of-two faster than the rate constant (Sander et al., 2006) used in WACCM4. Recently, the new JPL recommendation (Sander et al., 2011) was updated towards a higher value. However this remains an open issue and the possibility of increasing the agreement between observed data and model results by uploading the rate constant should be taken into account.

It is puzzling that ClO and ClONO<sub>2</sub> are quite well reproduced by WACCM4 in January 2005 but not in October–November 2003. Funke et al. (2011) argued that in 2003 the background ClO was underestimated by models, leading to modelled ClO changes smaller than observed. In contrast, in 2005 WACCM4 reproduces quite well the background ClO VMRs, and hence also the modelled SPE-induced changes are closer to observed data. A possible explanation could be that the chemical schemes of ClO formation/loss under twilight conditions are not well represented in the models. Another possible explanation could be that models tend to underestimate the actual isentropic mixing within the vortex. In 2003, a weaker mixing of “illuminated” and “dark” air masses inside the partially illuminated vortex could lead to lower modelled ClO abundances in the polar regions. In contrast, with a well established vortex coincident with the terminator as in 2005 the modelled results should be closer to the satellite data. The lower ClO abundance in the models during October 2003 is the main reason for the model underestimation of ClONO<sub>2</sub>. In contrast, in January 2005 the modelled ClO VMRs are similar to experimental values, hence also WACCM4 ClONO<sub>2</sub> output reproduces the MIPAS data well.

The SPE-induced HOCl changes recorded by MLS in January 2005 are similar to the HOCl variations recorded by MIPAS during the Halloween storm (about 0.25 ppbv). Even if we consider only the enhancement induced by the first SPE on 18 January (close to 0.2 ppbv at 1.5 hPa), therefore ignoring the additional increase due to the second SPE, the depletion still remains comparable. Because of the lower ionization of January 2005 (about 4–5 times less, see Sect. 5.1) this seems difficult to explain. The amount of SPE-induced HOCl enhancement is linked to the abundance of HO<sub>2</sub> that

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depends on the ionization; however also the seasonal conditions could play an important role. Since, under nighttime conditions,  $\text{HO}_x$  is in steady state also during a SPE, its abundance is proportional to ionization. The  $\text{HO}_x$  production induced by SPEs occurs mainly in the form of OH and H (Solomon et al., 1981). Then, OH can react with ozone and H, in presence of a third body, with oxygen molecules, both leading to  $\text{HO}_2$  formation. As the production of  $\text{HO}_2$  depends on the density, it follows that the  $\text{HO}_2/\text{OH}$  ratio decreases from the mesosphere down to the stratosphere under SPE conditions. Moreover, since the background of OH is higher than  $\text{HO}_2$  in the daytime upper stratosphere and mesosphere (i.e., the region affected by SPE), compared with its very low abundances at nighttime, this ratio is more elevated at nighttime. On the other hand,  $\text{HO}_x$  losses are dominated by reactions between OH and  $\text{HO}_2$ . In this way, the  $\text{HO}_2/\text{OH}$  ratio, i.e.  $\text{HO}_x$  partitioning, influences the HOCl formation since it controls the availability of  $\text{HO}_2$ . Therefore, the higher SPE-induced  $\text{HO}_2/\text{OH}$  ratio in January 2005, due to the almost complete nighttime conditions, contributes to an HOCl enhancement similar to that of 2003 despite the lower ionization and total  $\text{HO}_x$  production.

The recent studies of Jackman et al. (2011) and Funke et al. (2011) highlighted that models tend to overestimate the SPE-induced production of  $\text{HO}_2$  and  $\text{H}_2\text{O}_2$  and, hence, the  $\text{HO}_2/\text{OH}$  ratio under SPE conditions. As discussed, this ratio plays an important role in the partitioning of chlorine species. The MLS  $\text{HO}_x$  data could help us to constrain the  $\text{HO}_2/\text{OH}$  ratio. However the presence of a bias in MLS  $\text{HO}_2$  (Pickett et al., 2008), which normally requires taking day-night differences to be removed, prevents an accurate analysis (SPE-induced  $\text{HO}_2$  variations are evident mainly under nighttime conditions). The biases are typically small compared with SPE effects. Using differences of the nighttime data between two close time periods (to minimize the bias effect) it is possible to highlight SPE-induced  $\text{HO}_2$  changes (Jackman et al., 2011) although in this way the exact  $\text{HO}_2/\text{OH}$  ratio can be affected.

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## 8 Conclusions

For the first time the main contributors to the Cl<sub>y</sub> family recorded by two different satellites have been shown and compared with a state-of-the art chemistry climate model under SPE conditions. MLS measurements of chlorine species (i.e., HCl, ClO and HOCl), N<sub>2</sub>O, and temperature together with MIPAS ClONO<sub>2</sub> recorded during and after the SPEs of 17 and 20 January 2005 have been analysed and modelled with version 4 of the Whole Atmosphere Community Climate Model (WACCM4). SPE-induced chlorine repartitioning (HCl → ClO → HOCl and, after the events, ClONO<sub>2</sub>) has been identified. In the observations, HCl decrease is present at almost all the investigated altitudes. The strongest decrease (of less than 0.25 ppbv) occurs uniformly in a rather wide region ranging from 1.5 to 6 hPa on 21–22 January and it recovers in a week in the upper stratosphere. However below 4 hPa, this decrease is unlikely to be related to the SPE. Instead, it seems to be related to background variations.

ClO abundances did not undergo very large changes. However, a ClO decrease of up to 0.1 ppbv has been identified in the upper stratosphere at high latitudes, whereas a ClO increase of similar magnitude has been found close to the terminator. An HOCl enhancement of more than 0.25 ppbv occurred roughly in the same region where the ClO depletion took place. This confirms that HOCl is the main active chlorine species produced by proton forcing under nighttime conditions. Lower HOCl changes (~0.1 ppbv) have been detected at illuminated latitudes. There, SPE-induced active chlorine is present in the form of enhanced HOCl and ClO. An HOCl enhancement accompanied by a similar decrease in HCl was observed by MLS also during the SPEs of December 2006 (i.e., under similar seasonal conditions), further corroborating the hypothesis of chlorine activation under SPEs.

WACCM4 generally reproduces the SPE-induced variability in the chlorine species. Some differences in the geographic distribution of these changes arise from the underestimation of latitudinal redistribution above about 2 hPa. ClO decreases are reproduced quite well at high latitudes, whereas the ClO enhancements in the terminator

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region are underestimated. Moreover, contrary to our expectations, WACCM4 perturbed and base simulations indicate that this ClO increase is not due to SPEs. HOCl variations are modelled very well with respect to both magnitude and geographic distribution. WACCM4 also reproduces quite well the HCl depletion in the mesosphere (above  $\sim 1.5$  hPa) but underestimates or does not show the observed decrease between 1.5 and 4 hPa. ClONO<sub>2</sub> changes observed by MIPAS have also been examined. An increase of 0.2 ppbv around 1 hPa in the week after the SPE occurrence has been shown. WACCM4 simulations with and without proton forcing indicate that this increase is dominated by background variability, although SPE-induced production might contribute by 0.1 ppbv.

The most striking differences between experimental data and modelling results are (i) the underestimation of the HCl depletion mainly below 1.5 hPa and (ii) the lack of enhanced ClO in the terminator region in the WCCM4 simulations. The possible background variation highlighted in Figs. 4 and 5 could contribute to explain the HCl discrepancies. On the other hand, more accurate simulations of HO<sub>x</sub> under SPE conditions (Funke et al., 2011; Jackman et al., 2011) could also clarify the ClO differences. In general, WACCM4 simulations reproduce better the impact of SPEs on chlorine species for the January 2005 SPE than for the 2003 Halloween storm (Funke et al., 2011). The major reason for this difference is that the vortex was well established, coincident with the terminator, in January 2005, while the early winter vortex in October 2003 was partially illuminated. However, also during January 2005, the WACCM4 simulations did not fully capture all the measured changes and future work will be necessary to understand and remove remaining minor model-data inconsistencies.

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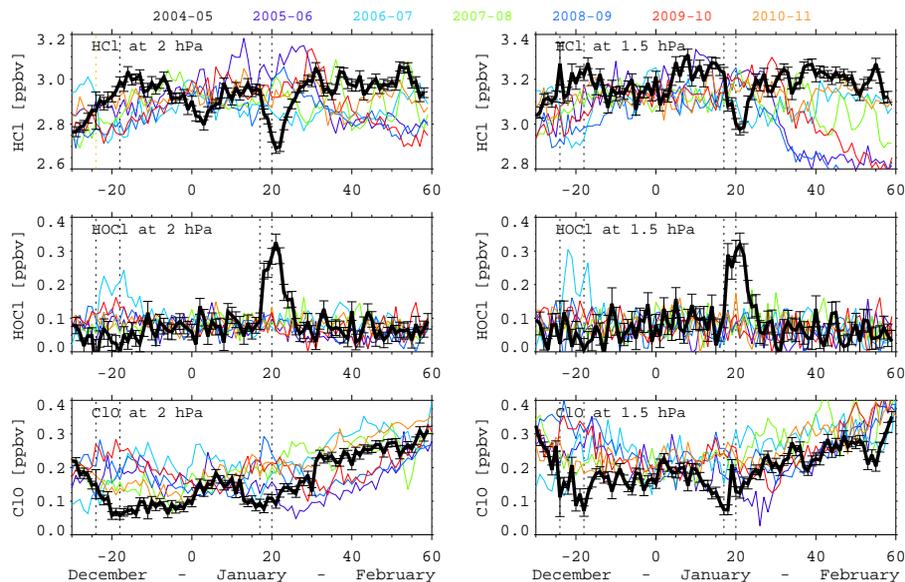
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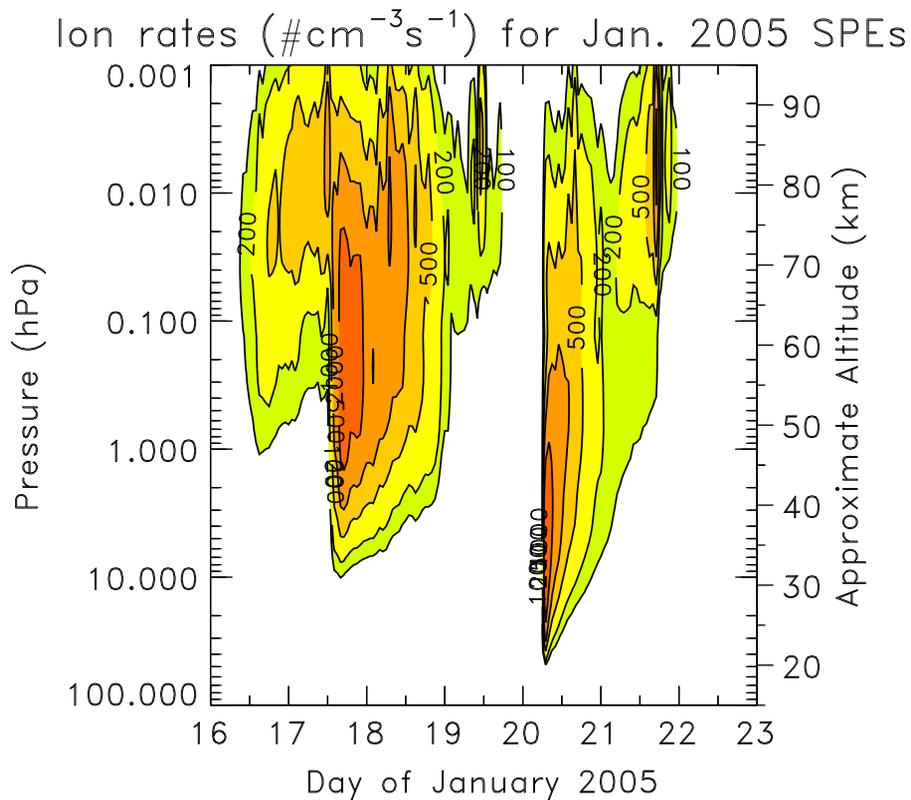
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**Fig. 1.** From top to bottom: time series of daily MLS HCl, HOCl and ClO at 2 hPa (left column) and 1.5 hPa (right column) averaged over 70–82° N from December to February for 2004–2011 years. The vertical dashed lines show the SPEs occurrence. The estimated precision (vertical error bars) for 2004–2005 is reported as a reference.

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**Fig. 2.** Hourly average ion pair production rates ( $\text{cm}^{-3}\text{s}^{-1}$ ) as a function of time for 16–22 January 2005. The contour levels are 100, 200, 500, 1000, and 2000  $\text{cm}^{-3}\text{s}^{-1}$ . Proton flux data from GOES-11 were used to compute ion pair production profiles.

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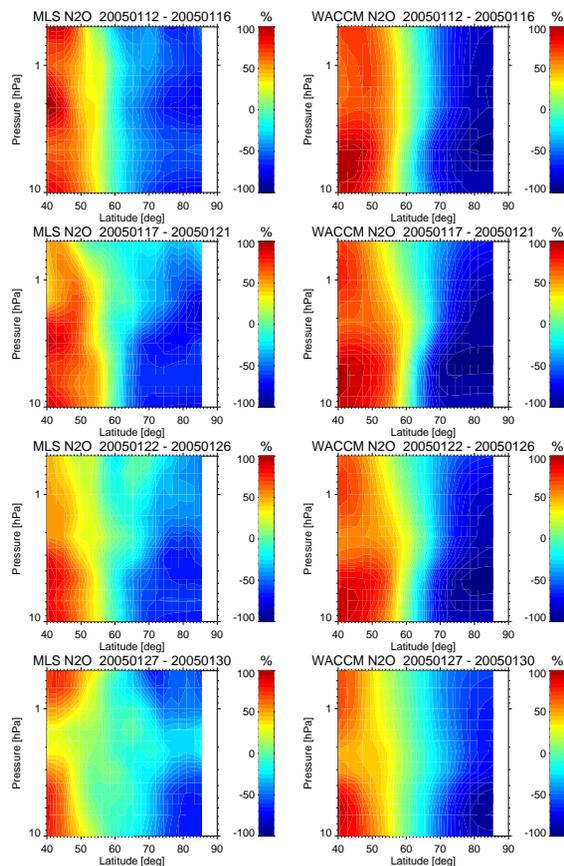
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**Fig. 3.** Zonal mean relative anomaly of  $N_2O$ , with respect to the average ( $40\text{--}82^\circ N$ ), expressed in percent. From top to bottom: zonal mean anomaly of  $N_2O$  before (12–16 January), during (17–21 January) and after (22–26 January and 27–30 January) the SPEs of 17 and 20 January 2005. Left panels: MLS data. Right panels: WACCM4 result.

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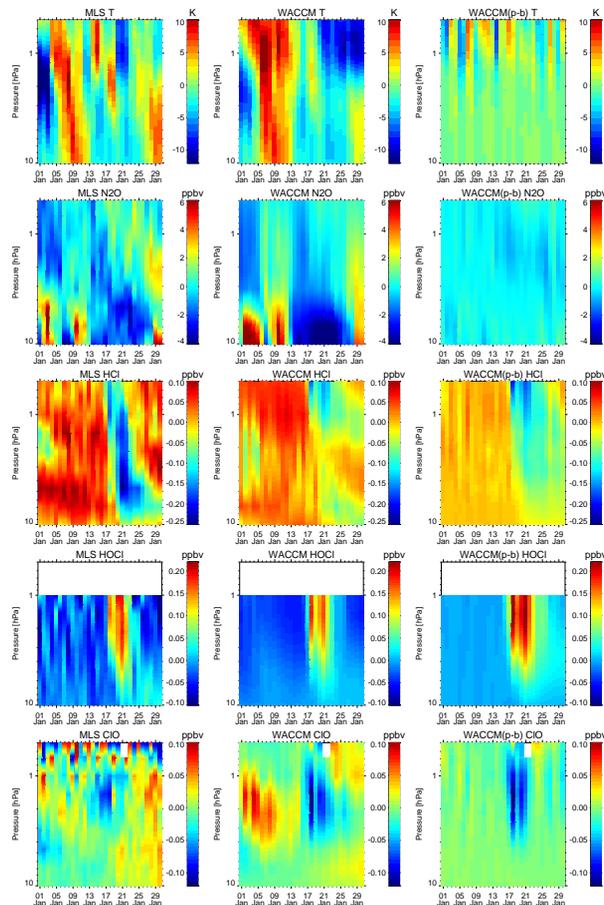
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**Fig. 4.** From top to bottom: temporal development of daily MLS T, N<sub>2</sub>O, HCl, HOCl, ClO during January 2005, averaged over 70–82° N, relative to background conditions (January average). Left panels: MLS data. Middle panels: WACCM4 result. Right panels: the model SPE response (perturbed – base).

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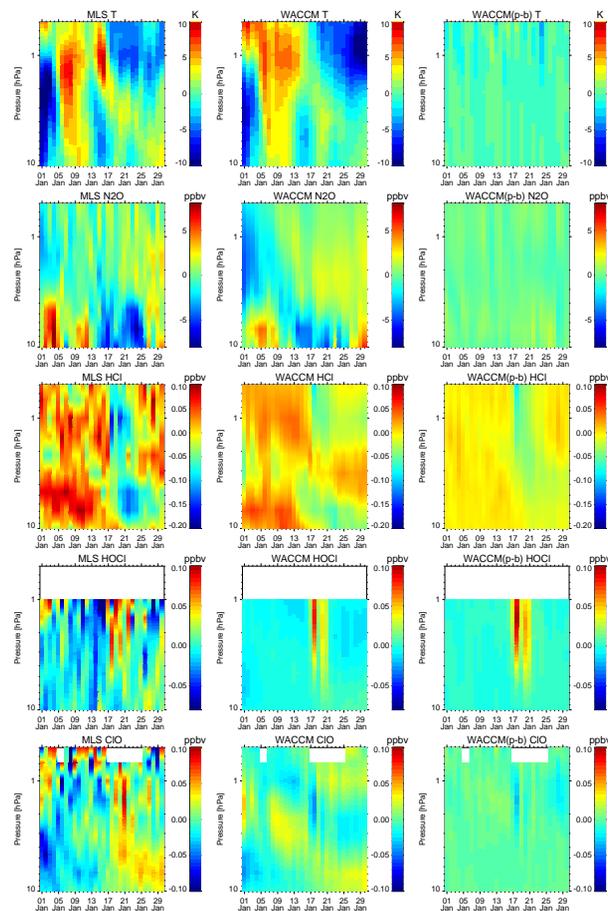
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**Fig. 5.** From top to bottom: temporal development of daily MLS T, N<sub>2</sub>O, HCl, HOCl, ClO during January 2005, averaged over 58–70° N, relative to background conditions (January average). Left panels: MLS data. Middle panels: WACCM4 result. Right panels: the model SPE response (perturbed – base).

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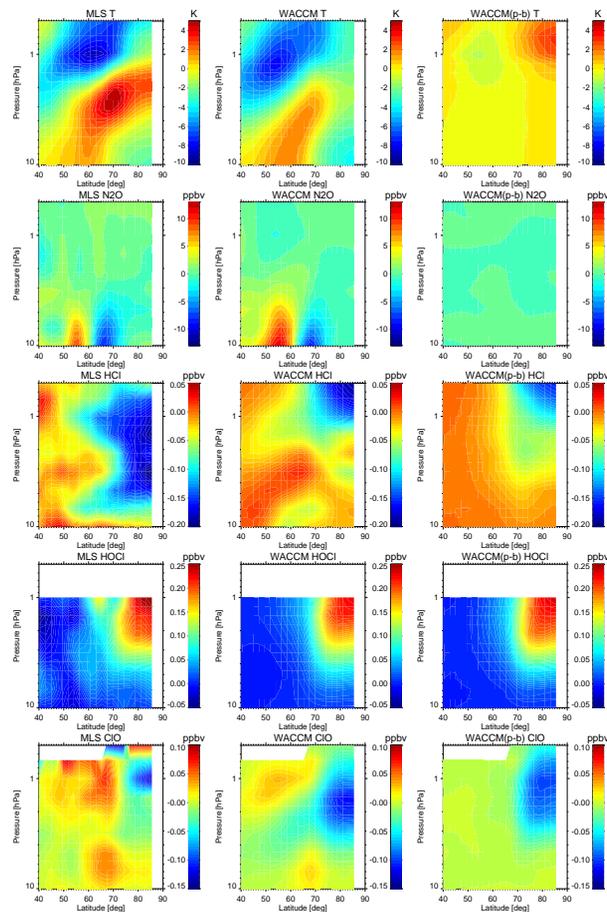
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**Fig. 6.** From top to bottom: zonal mean change (17–21 January 2005 with respect to 12–16 January 2005) of MLS T, N<sub>2</sub>O, HCl, HOCl, ClO at 40–90°N. Left panels: MLS data. Middle panels: WACCM4 result. Right panels: the model SPE response (perturbed – base).

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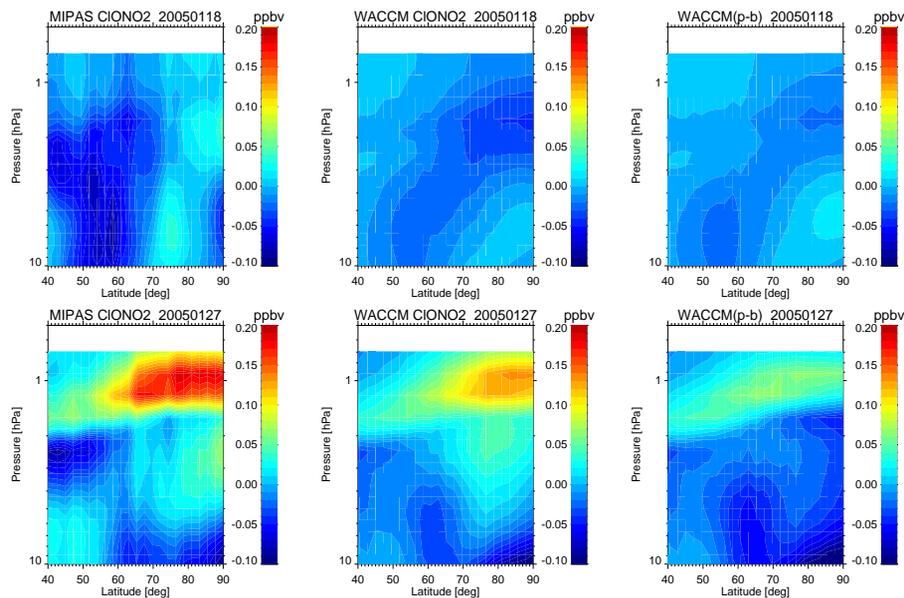
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**Fig. 7.** CIONO<sub>2</sub> zonal mean change on 18 (upper panels) and 27 January 2005 (lower panels) with respect to 16–17 January 2005 at 40–90° N. Left panels: MIPAS data. Middle panels: WACCM4 result. Right panels: the model SPE response (perturbed – base).

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